

A time-dependent radiative model for the atmosphere of the eccentric transiting planets

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Abstract. We present a time-dependent radiative model for the atmosphere of the transiting planets that takes into account the eccentricity of their orbit. We investigate the temporal temperature and flux variations due to the planet-star distance variability. We will also discuss observational aspects with *Spitzer* measurements.

1. Introduction

Why taking the eccentricity into account? The major theoretical models of atmospheric structure of exoplanets consider the planet only into two conditions:

- either the flux received by the planet coming from the star is averaged over an hemisphere of the planet in a static case ;
- or some sort of redistribution day to night of the incoming flux is included.

In none of these cases the fact that the flux actually varies with time is taken into account. Even in the case of less eccentric planet, like WASP-10 b with a low eccentricity of 0.057 (Christian *et al.* 2008) this represents a variation $\sim 25\%$ in the received flux between periastron and apoastron.

We focused our study on two planets: HD 17156b, the most eccentric transiting planet, and even if HD 80606b is not transiting, we choose this planet with the largest eccentricity to show the extreme case. Table 1 summarize the parameters of these two systems.

2. Model description

In order to model the planets atmosphere, we use a line-by-line radiative transfer code which is described in details in Iro *et al.* (2005).

The equation we have to solve is:

$$\frac{dT}{dt} = \underbrace{-\frac{mg}{C_p} \frac{dF_\star}{dp}}_{\text{Heating rate}} - \underbrace{\left(-\frac{mg}{C_p} \frac{dF_{\text{IR}}}{dp}\right)}_{\text{Cooling rate}} \quad (2.1)$$

where T is the temperature, t time and we divided the flux F into 2 quantities: the thermal flux emitted by the planet (upward-downward) F_{IR} and the stellar net flux (downward-upward) F_\star , so that we have $F = F_{\text{IR}} - F_\star$.

We take advantage of the fact that the time is explicitly included in the equation by introducing a variation of the insolation at each time step. This variation can be

Table 1. Planetary and stellar parameters

	Parameters	HD 17156b	HD80606 b
Orbit	e	0.67	0.93
	a [AU]	0.159	0.432
	d_{\min} [AU]	0.052	0.030
	d_{\max} [AU]	0.266	0.834
	P [days]	21.22	111.78
	ω [°]	121	300
Planet	Radius [R_J]	0.96	1.1 ^a
	Mass [M_J]	3.11	5 ^b
	P_{spin} [days]	3.8	1.93
Star	Type	G0V	G5
	Radius [R_{\odot}]	1.47	1.05
	[Fe/H]	0.24	0.43
Refs.		Gillon <i>et al.</i> (2007)	Langton & Laughlin (2008)

Notes:

^a This planet is not transiting. The value of the radius has been predicted by (4) ;

^b We adopted this value of the mass (inclination unknown).

due to the non-constant planet-to-star distance during the eccentric orbit and/or to the rotation of the atmosphere since the eccentric planets have a remnant spin. They are indeed supposed to be in pseudo-synchronous rotation. We calculated the rotation period using the prescription of Hut (1981) (see Table 1).

Method

Here is the scheme of our calculations:

Static calculations

The first thing we do is calculate the thermal structure of a planet with the planet located at the semi major axis.

We do so iteratively. That means that

- (a) for an initial thermal profile we calculate the heating and cooling rates
- (b) then make them converge by modifying the thermal profile according to Eq. 2.1 setting the left-hand term to 0 (*thermal equilibrium*)
- (c) if the profile is different by the end of process (b) ; we begin process (a) with the new profile ; if it is the same that means that we have reached the static solution.

Orbit calculation

We then calculate for each time the actual planet-star distance and take the distance factor into account to correct for the right insolation. This is a modified version of point (b) from above.

We obtain a thermal profile as well as the spectra for each time step. The results of this step is shown in Fig. 1 and Fig. 2 for the thermal structures and in Fig. 3 for the fluxes.

Rotation calculation

Eventually we add the rotation of the planet. These eccentric planets are indeed supposed to not be completely synchronized. But we choose to implement the rotation along with a multi-latitude / multi-longitude approach (and a parallelization of the code). This step is still ongoing (see the Perspectives section for further details).

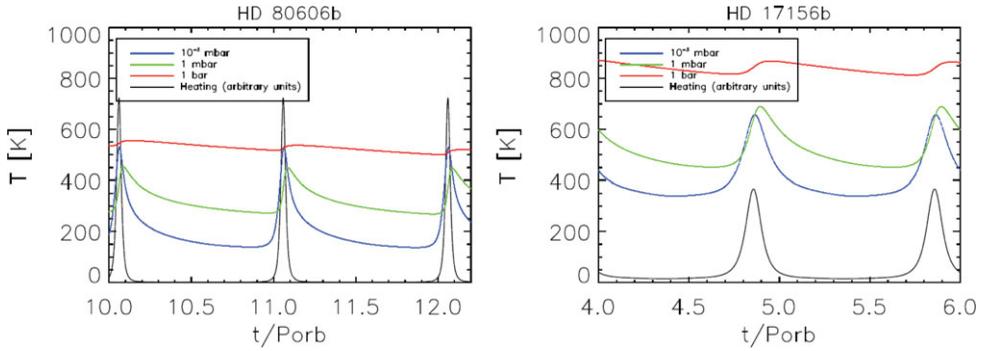


Figure 1. Temperature of the planet atmospheres at selected pressure levels as a function of time for the condition where the flux received by the planet is averaged over its sphere. We let the calculation make several orbital revolutions in order to reach a periodic solutions (note the different scales).

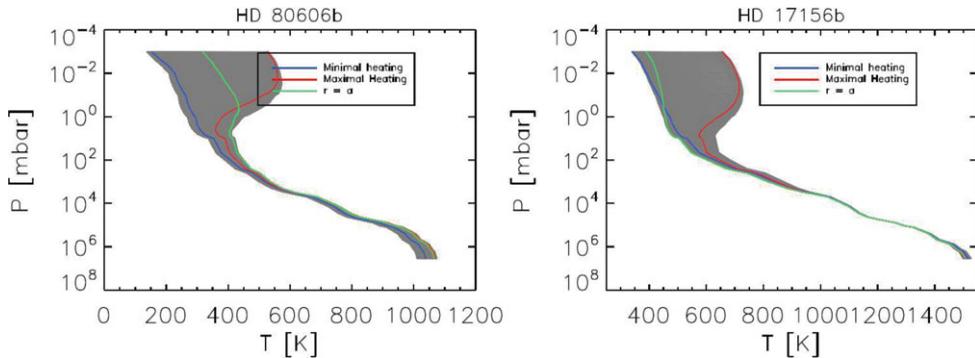


Figure 2. Temperature Structures of the atmospheres at selected orbit positions (*lines*) and for every times (*shaded area*) (note the different scales).

2.1. Limitations

Here are some simplifications that we have done:

- The chemical composition is not changed as the thermal profile does. The species are calculated at the semi major axis distance.
- In the same way, the heating rates and cooling rates are for now only calculated at the semi major axis. Then only the varying *stellar heating* is accounted for.
- we have not included any clouds due to the uncertainties on the formation processes.

Discussion

Fig. 3 compares the ratio of the planet flux over the star flux as a function of wavelength. It is worth noting that before comparing directly with *Spitzer* observations, we have to take into account the phase effects. Here, we have an average over the planet. During secondary eclipses we observe the illuminated hemisphere, this leads to very different values of the flux. In the next development of the code (see the Perspective section), this will be investigated.

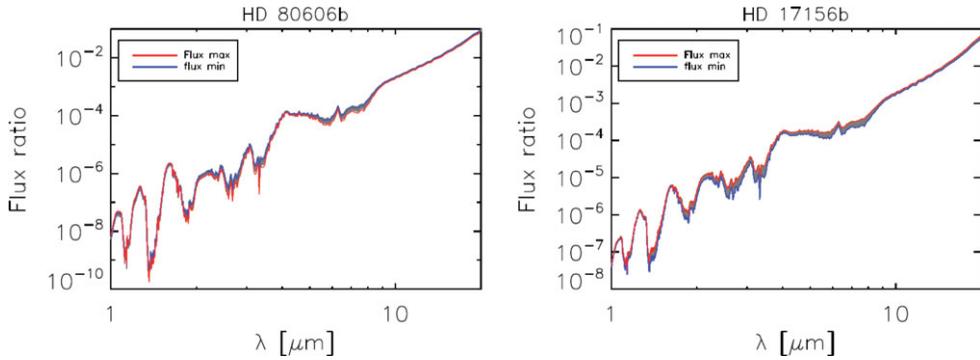


Figure 3. Ratio of the planet flux to the star flux as a function of wavelength at two positions (*lines*) and for every times (*shaded area*).

3. Perspectives

- We have benefited from the parallelization capabilities of our code to include a different heating pattern depending on the longitude and latitude. By adding also the planet rotation with the same method as in Iro *et al.* (2005), we will be able to construct a thermal map of the planet as a function of time. By integrating the view corresponding to the observation, we will also be able to calculate a full light curve in any wavelength band.

- We will then also include a calculation of the heating rates and cooling rates for each planet as well as the corresponding thermal profiles and chemical composition.

- Finally, we will also calculate the above quantity as a function of time. This will require a huge increase of the calculation speed. The use of the correlated k-coefficient method (opposed to line-by-line) will probably be necessary.

4. Conclusion

We presented a time-dependent radiative model, still under development, for the atmosphere of the transiting planets. This model takes into account the eccentricity of their orbit and in its next version will be able to completely map the planet, including the rotation effects on the insolation.

The preliminary results already show that an eccentric planet under heating conditions is very different to one on a circular orbit, even for a low eccentricity, although several steps need to be completed before reaching a fully self-consistent model.

References

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