

CS(3-2) in Nearby Starburst Galaxies M82 and NGC253

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Dense gas in molecular cloud cores is essential in researches of extragalactic star formation (e.g. Solomon, Downes, and Radford^[1]). To determine the physical relations between star-forming regions and dense gas, we have observed CS ($J=3\rightarrow 2$) in starburst galaxies M 82 and NGC253. The $J = 3$ level of CS is 14.1 K and the critical density for excitation is $4 \times 10^5 \text{ cm}^{-3}$.

Maps of the CS ($J=3\rightarrow 2$) of NGC253 are illustrated in Fig. 1. The fraction of continuum emission is so small that it has not been subtracted in Fig. 1. The CS ($J=3\rightarrow 2$) emission is separated in velocity; the red- and blue-velocity components are located west and east of the nucleus, respectively. This is similar to the HCN ($J=1\rightarrow 0$) and HCO⁺ ($J=1\rightarrow 0$) maps by Carlstrom *et al.*^[2].

A map of the 2 mm continuum and maps of the CS ($J=3\rightarrow 2$) of M82 are illustrated in Fig. 2. The continuum has been subtracted at each visibility spectrum before the CS ($J=3\rightarrow 2$) maps are made. The CS ($J=3\rightarrow 2$) emission is distinctly stronger in the east side of the center than in the west side. This is quite different from any other tracer; centimeter-^[3] and millimeter-wave radio continuum (Figure 2 and ref. [2]), mid-infrared continuum^[4], HI^[5], CO ($J=1\rightarrow 0$)^[6], CO ($J=2\rightarrow 1$)^[7], HCN ($J=1\rightarrow 0$)^[2], and HCO⁺ ($J=1\rightarrow 0$)^[2] show two peaks with similar strengths (or rather stronger in the west). The upper limit at the west side is consistent with the CS ($J=3\rightarrow 2$) spectrum by Mauersberger *et al.*^[8]. The peculiar distribution of CS ($J=3\rightarrow 2$) in M82 is not explained simply by the difference between the temperatures of the dense gas in both sides of the center since the multi-line study by Wild *et al.*^[9] indicates that the line strengths of higher transitions of CO, C¹⁸O, HCN, and HCO⁺ are not weak at the west peak. The underlying mechanism for this difference is not yet understood.

REFERENCES

- [1] Solomon, P.M., Downes, D., and Radford, S.J.E. 1992, *ApJL*, **387**, L55
- [2] Carlstrom, J.E. 1988, Ph.D Thesis at U. C. Berkeley
- [3] Kronberg, P.P., Biermann, P., and Schwab, F.R. 1985, *ApJ*, **291**, 693
- [4] Rieke, G.H. *et al.* 1980, *ApJ*, **238**, 24
- [5] Weliachew, L., Fomalont, E.B., and Greisen, E.W. 1984, *A&A* **137**, 335
- [6] Lo, K.Y. *et al.* 1987, *ApJ*, **312**, 574
- [7] Loiseau, N. *et al.* 1990, *A&A*, **228**, 331
- [8] Mauersberger, R. *et al.* 1991, *A&A*, **247**, 307
- [9] Wild, W. *et al.* 1992, *A&A*, **265**, 447

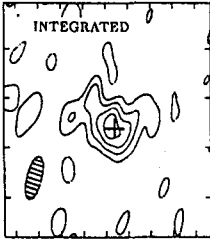
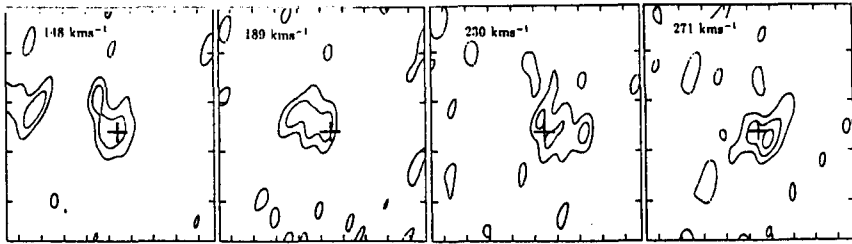


Figure 1 CS ($J=3\rightarrow 2$) maps of NGC 253. The separation between adjacent tick marks on the declination axis is $15''$.

Figure 2 CS ($J=3\rightarrow 2$) maps and a 2 mm continuum map of M 82. The separation between adjacent tick marks on the declination axis is $5''$.

