

Turbulent magnetic pressure instability in stratified turbulence

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Abstract. A reduction of total mean turbulent pressure due to the presence of magnetic fields was previously shown to be a measurable effect in direct numerical simulations. However, in the studied parameter regime the formation of large-scale structures, as anticipated from earlier mean-field simulations, was not found. An analysis of the relevant mean-field parameter dependency and the parameter domain of interest is conducted in order to clarify this apparent discrepancy.

Keywords. turbulence, MHD, sunspots

Strong magnetic fields at the solar surface are generally thought to originate from the coherent rise of magnetic flux tubes from the tachocline through the solar convection zone. While the idea is elegant, the question remains whether it can be considered as more than a toy model, as the physics of the creation and rise of these flux tubes is not sufficiently understood Parker, 2009. Thus far any ‘successful’ numerical simulation of this process had to rely on strong assumptions, be it in the initial conditions or in simplified equations such as the thin flux tube approximation Spruit, 1981. As such it makes sense to explore alternative mechanisms of magnetic structure formation. Several models have been proposed where large-scale magnetic field concentrations are created through instabilities at the solar surface (Kitchatinov & Mazur, 2000, Brandenburg *et al.* 2010a, hereafter BKR).

Turbulence is generally associated with enhanced transport effects. However, it can also generate structures on much larger scales than its driving scales; see as an example the inverse cascade in 2D hydrodynamic turbulence (Kraichnan, 1967) or 3D MHD turbulence with magnetic helicity (Frisch *et al.* 1975). We study here the interaction of the turbulence with a background magnetic field. From the approximate conservation of total turbulent energy E_{tot} , BKR find a reversed feedback from the magnetic fluctuations on the turbulent pressure (Rogachevskii & Kleeorin, 2007, hereafter RK):

$$P_{\text{turb}} = -\frac{1}{6}\bar{b^2}/\mu_0 + \frac{2}{3}E_{\text{tot}}.$$

It can be seen that the effective mean magnetic pressure force is reduced and can be reversed in a certain parameter range. It was suggested by RK that this positive feedback could lead to an instability, resulting in the concentration of magnetic flux.

BKR confirmed the validity of approximate turbulent energy conservation using direct numerical simulations (DNS) of homogeneous isothermal turbulence and they also demonstrated the basic phenomenon of magnetic flux concentration through the interaction between turbulence and the mean Lorentz force in mean-field MHD simulations which led to a linear instability for sufficiently strong stratification. These results were recently corroborated using DNS of inhomogeneous (stratified) isothermal turbulence of

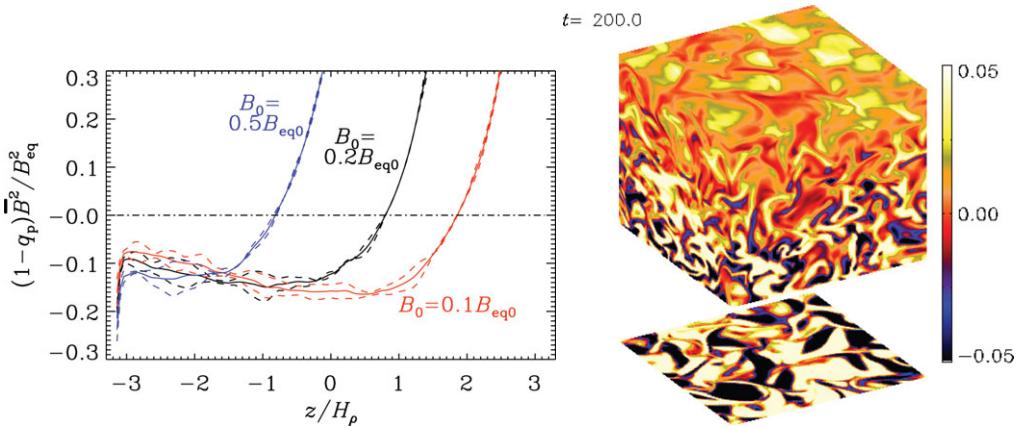


Figure 1. *Left:* Normalized effective mean magnetic pressure as a function of depth for $B_0 = 0.1B_{eq0}$, $B_0 = 0.2B_{eq0}$, and $B_0 = 0.5B_{eq0}$ using $Re = 120$ and $Pm = 1$. Adapted from Brandenburg *et al.* 2010b. *Right:* Visualization of $B_y - B_0$ on the periphery of the computational domain $B_0 = 0.1B_{eq0}$, and $Pm = 2$. Adapted from Kemel *et al.* 2010.

a cubic computational domain. An example is shown in Figure 1 where we show that the normalized effective mean magnetic pressure does indeed become negative in a large part of the domain for different imposed field strengths. However, the formation of large-scale structures is not observed, as can be seen from the right-hand panel of Figure 1.

It is important to understand whether there is really a conflict between DNS and mean-field results, or if these two studies simply apply to different parameter regimes. In order to have a chance to resolve this question, we conduct a systematic parameter survey of the instability in the mean-field model. More specifically, we determine the functional dependence of the growth rate on the input variables in order to find the relevant parameter space for the instability to develop.

In an isothermal stratified box we solve the equations in two dimensions:

$$\frac{\partial \bar{\mathbf{U}}}{\partial t} = -\bar{\mathbf{U}} \cdot \nabla \bar{\mathbf{U}} - c_s^2 \nabla \ln \bar{\rho} + \mathbf{g} + \bar{\mathcal{F}}_M + \bar{\mathcal{F}}_{K,\text{tot}}, \quad (0.1)$$

where

$$\bar{\rho} \bar{\mathcal{F}}_M = -\frac{1}{2} \nabla [(1 - q_p) \bar{\mathbf{B}}^2] + \bar{\mathbf{B}} \cdot \nabla [(1 - q_s) \bar{\mathbf{B}}] \quad (0.2)$$

is the mean-field Lorentz force and

$$\bar{\mathcal{F}}_{K,\text{tot}} = (\nu_t + \nu) (\nabla^2 \bar{\mathbf{U}} + \nabla \nabla \cdot \bar{\mathbf{U}} + 2\bar{\mathbf{S}} \nabla \ln \bar{\rho}) \quad (0.3)$$

is the total (turbulent and microscopic) viscous force with \mathbf{S} being the viscous stress tensor. In addition, we solve the continuity and uncurled induction equations,

$$\frac{\partial \ln \bar{\rho}}{\partial t} = -\bar{\mathbf{U}} \cdot \nabla \ln \bar{\rho} - \nabla \cdot \bar{\mathbf{U}}, \quad \frac{\partial \bar{\mathbf{A}}}{\partial t} = \bar{\mathbf{U}} \times \bar{\mathbf{B}} - (\eta_t + \eta) \bar{\mathbf{J}}. \quad (0.4)$$

We adopt a Cartesian coordinate system, (x, y, z) . The mean field is given by $\bar{\mathbf{B}} = (0, B_0, 0) + \nabla \times \bar{\mathbf{A}}$, and the vertical gravitational acceleration is $\mathbf{g} = (0, 0, -g)$. The other input parameters of the simulations are the sound speed c_s , the density at the top of the box ρ_{top} , the pressure coefficient q_p (RK), the magnetic Prandtl number Pm and the molecular diffusivity η .

We start by isolating the dependence of the growth rate on the turbulent diffusivity and find approximately linear behaviour. From a data survey we observe, depending on

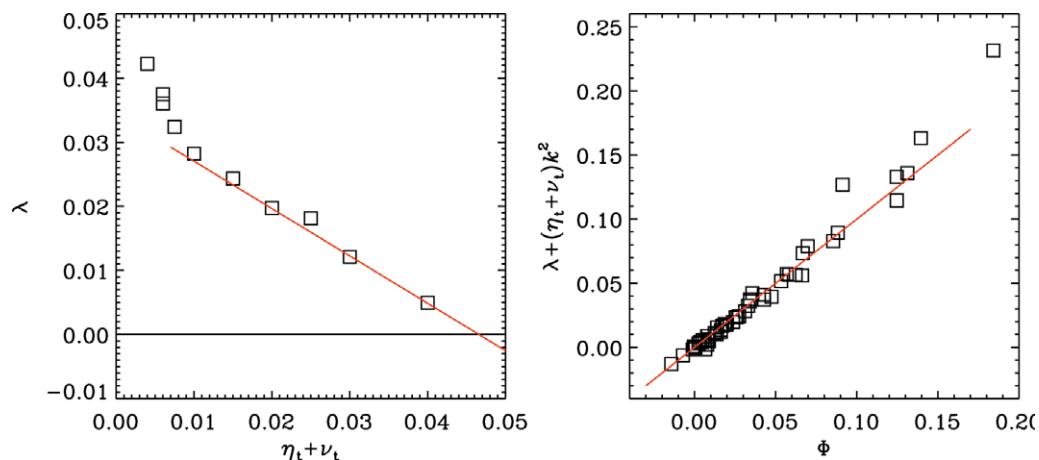


Figure 2. Left: dependence of growth rate on turbulent diffusivity and viscosity. Right: modified growth rate as a function of Φ .

the applied field, different diffusivity values above which the instability does not develop, which would imply

$$\lambda = \Phi(B) - (\eta_t + \nu_t)k^2 \quad (0.5)$$

where $1/k$ is a length scale introduced for dimensional reasons. We find that Φ is proportional to the Alfvén speed at the top of the box, where the instability initiates. Eventually we arrive at the fit formula

$$\lambda + \left(1 + \text{Pr}_M^{\text{turb}}\right) \eta_t k^2 = v_A k_x \left(1 + q_p/q_p^*\right) \exp(\ell_z/H_\rho), \quad (0.6)$$

as shown in Figure 2. Here, $\text{Pr}_M^{\text{turb}} = \nu_t/\eta_t$, ℓ_z is some typical vertical length scale, and $H_\rho = c_s^2/g$ is the density scale height. The instability resulting from this feedback effect was verified but not observed to generate large-scale structures in DNS with the current scale separation and parameter range studied so far.

An extension of this work will be the inclusion of more physics (e.g. radiative cooling) in the DNS and the comparison with mean-field models in a 3D setup. As was already pointed out in BKR, the mean-field model predicts additional short-wavelength perturbations along the direction of the mean magnetic field, that were not included in the present study.

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