

## **Interpretation of Rotation Curves of Giant Galaxies in the Local Group with the Truncated Disk Model: the Milky Way**

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### **1. Introduction**

In some external edge-on spiral galaxies, such as NGC 891, 4244, 4565 and 5907, the observed light distribution exhibits a sharp truncation in the outer parts (e.g., Casertano 1983, hereafter C83). As a rule, such galaxies are known to be warped in their gas disk, the deformation starting at about their optical cut-off radius. These facts suggest a truncation in the disk density distribution. Moreover, the rotation velocity of these galaxies starts decreasing at about the same radius, which can also be considered as the “signature” of truncation (see C83).

Both giant spiral galaxies in the Local Group, the Milky Way and M31, have a dip in the average rotation velocity which can be produced by a truncation of the galactic disk. Ninković & Petrovskaya (1992) constructed from Haud’s (1981) rotation curve for M31 a model including, like C83’s for NGC 5907, a truncated exponential disk and a halo. The truncation distance was found to be close to the radius beyond which warping of the hydrogen layer of M31 begins.

For our Galaxy a preliminary mass distribution model was constructed which reproduces a dip in rotation velocity at about the solar radius  $R_0$  by Petrovskaya (1992). In this paper we present results of a more comprehensive study for the Galactic rotation curve constructed in Nikiforov (1999) (hereafter N99) with  $R_0 = 8.2$  kpc and the solar circular velocity of  $216.5 \text{ km s}^{-1}$ .

### **2. Rotation curve of the Milky Way**

We use the recent data on the rotation of the neutral hydrogen from the whole 21-cm line profile and the tangent points, and of molecular clouds with CO emission (for details see N99). Our rotation curve of the Milky Way (Figure 1 in N99 and Figure 1 in this paper) has a significant dip, on average, of  $\sim 20 \text{ km s}^{-1}$  between relative radii of  $x \equiv R/R_0 = 0.75$  and 1.15. At larger distances the rotation velocity  $\theta$  initially increases and then remains virtually constant at  $x > 1.7$ –1.8.

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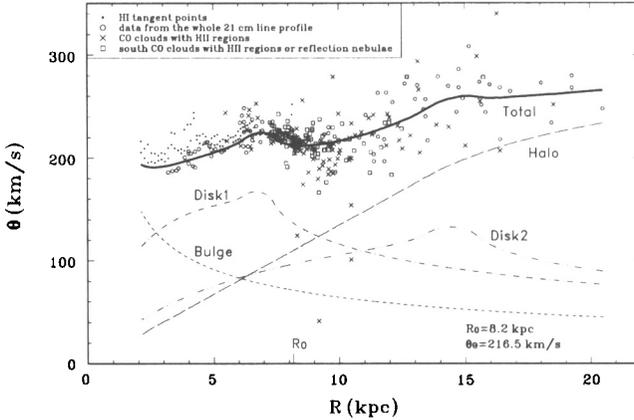


Figure 1. A mass model of the Milky Way fitted to the rotation curve from HI and CO data.

It is usual practice to ignore these details of the rotation curve in constructing a mass distribution model for our Galaxy (e.g., Merrifield 1992). On the other hand, Lépine & Amaral’s (1995) attempt to reproduce the minimum at  $x \sim 1.1$  exclusively by adding a perturbation from spiral arms fared poorly. Besides, it is well known that the Milky Way also has a warp of the HI layer for  $x > 1.1$  (e.g., Gerasimov & Petrovskaya 1990), the thickness of the layer starting to rise steeply at  $x > 0.8-0.9$ . On this basis we suppose that the dip in density of the Galactic disk at  $R < R_0$  could, at least in part, explain the decrease in  $\theta$  near  $R_0$ .

### 3. Model

We constructed a four-component model for the Milky Way consisting of two truncated exponential disks (inner and extended), a bulge and a “dark halo”. The inner disk is required to reproduce the dip in  $\theta$  at  $x < 1$  by a fall in disk density. The extended disk is called for a representation of the disk matter in the vicinity of the Sun and in the outer Galaxy where the disk component has a cut-off at  $x \sim 1.7$  (Porcel et al. 1997 and references therein). The “break” in the rotation curve at  $x = 1.7-1.8$  (Figure 1) may be associated with the latter detail.

We use C83’s expressions for density distribution in the “softly truncated disk”:

$$\rho(R, z) = \rho_0(R)[\text{ch}(z/z_0)]^{-2}, \tag{1}$$

$$\rho_0(R) = \begin{cases} \rho_{00} \exp(-R/h), & R \leq R_d \\ \rho_0(R_d)[1 - (R - R_d)/\delta], & R_d \leq R \leq R_d + \delta \\ 0, & R \geq R_d + \delta, \end{cases} \tag{2}$$

where  $R$  is the distance from the Galactic axis,  $z$  is the distance from the Galactic plane,  $z_0$  is the disk half-thickness ( $=100$  pc, i.e. the mean thickness of HI and CO layers),  $R_d$  is the cut-off radius,  $\delta$  is the cut-off length. The central surface density  $\mu_0$  equals  $2\rho_{00}z_0$ .

Note that we do not suggest any physical difference between the two disks; this formal decomposition is taken only to represent dips in disk density.

As the bulge model, we use a truncated Hubble law spheroid:

$$\rho(q) = \begin{cases} \rho(0) [1 + (q/q_b)^2]^{-3/2}, & q \leq a_b \\ 0, & q > a_b, \end{cases} \tag{3}$$

$$q = [R^2 + (z/\varepsilon)^2]^{1/2}, \quad \varepsilon = 0.6. \tag{4}$$

This component has only marginal dynamical significance in the range of distances probed by our rotation curve.

The density distribution used for the halo is constant in the inner part and pseudo-isothermal in the outer part:

$$\rho(q) = \begin{cases} A \left[ \frac{1}{1 + (q/q_h)^2} - \frac{1}{1 + (a_h/q_h)^2} \right], & R_h \leq q \leq a_h \\ \rho(R_h), & q \leq R_h \\ 0, & q \geq a_h, \end{cases} \tag{5}$$

$$q = (R^2 + z^2)^{1/2}. \tag{6}$$

This complication is based on our modelling results: we found the scale parameter  $q_h$  to increase without limit during the fitting procedure for the pure pseudo-isothermal model.

#### 4. Model Fitting and Results

By assuming values of model parameters ( $q_b$ ,  $q_h$  and  $a_h$ ), which cannot be found uniquely, we derived others by a fitting to the rotation data. The best-fit solution, which minimizes the quantity  $\sum_j (\theta_{\text{obs}} - \theta_{\text{mod}})_j^2 / \sigma_{\theta_j}$ , was found by quasi-Newton algorithm for finding minimum of a function of more than one variables. The mass of the extended disk was normalized to give a local surface density of  $50 M_\odot \text{pc}^{-2}$  (Kuijken & Gilmore 1989, Holmberg & Flynn 2000). Results are given in Table 1. The mass of the Galactic disk is  $5.85 \times 10^{10} M_\odot$ , the total mass of the Galaxy in the central 20.5 kpc of this model is  $(3.3 \pm 0.4) \times 10^{11} M_\odot$ .

The mass model is presented in Figure 1. This model reproduces the dip in  $\theta$  of  $13 \text{ km s}^{-1}$  at  $R \sim R_0$  and the ‘‘break’’ at  $x = 1.7\text{--}1.8$ . The model dip is in good agreement with the observed one from the HI ( $\approx 10 \text{ km s}^{-1}$ ). The observed greater dip of  $30 \text{ km s}^{-1}$  from molecular clouds is most probably induced not by a drop in the disk density but by a perturbation from a spiral wave.

#### 5. Conclusion

Rotation curves of both giant spiral galaxies in the Local Group can be reproduced in detail with the truncated exponential disk model. In the case of the

Table 1. Mass model solution for the Milky Way galaxy.

Component	Parameters	Local Density	Mass
Inner Disk	$h = 2.7$ kpc $R_d = 6.21$ kpc $R_d + \delta = 7.4$ kpc	$\mu_0 = 808 M_\odot \text{pc}^{-2}$ $\rho_{00} = 4.04 M_\odot \text{pc}^{-3c}$	$0$ $2.68 \times 10^{10} M_\odot^c$
Extended Disk	$h = 8$ kpc $R_d = 13.5$ kpc $R_d + \delta = 15.5$ kpc	$\mu_0 = 145 M_\odot \text{pc}^{-2}$ $\rho_{00} = 0.724 M_\odot \text{pc}^{-3c}$	$\mu_\odot = 50 M_\odot \text{pc}^{-2a}$ $3.17 \times 10^{10} M_\odot^c$
Bulge	$q_b = 0.4$ kpc <sup>d</sup> $a_b = 2.1$ kpc	$\rho_0 = 14.4 M_\odot \text{pc}^{-3}$	$0$ $0.95 \times 10^{10} M_\odot^c$
Halo	$q_h = 2.5$ kpc <sup>d</sup> $a_h = 70$ kpc <sup>b</sup> $R_h = 13.4$ kpc	$A = 0.31 M_\odot \text{pc}^{-3}$ $\rho(R_h) = 0.01 M_\odot \text{pc}^{-3c}$	$\rho_\odot = 0.01 M_\odot \text{pc}^{-3c}$ $2.60 \times 10^{11} M_\odot^c$ ( $R \leq 20.5$ kpc)

<sup>a</sup>Kuijken & Gilmore (1989), Holmberg & Flynn (2000)

<sup>b</sup>Ninković (1992)

<sup>c</sup>Calculated from found and accepted parameters

<sup>d</sup>Accepted after a tentative fitting

Milky Way, there is evidence for two density dips in the disk which leave their “signatures” on the rotation curve. Conceivably the generation of these truncations might be induced by interactions of giant galaxies with small satellites or by rare encounters with more massive but presently remote LG member. This may be considered as a special problem of the LG dynamics, the solution of which requires data on proper motions.

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## References

- Casertano, S. 1983, MNRAS, 203, 735  
 Gerasimov, A. G., & Petrovskaya, I. V. 1990, Kinematika i Fizika Nebesnykh Tel, 8, 17  
 Haud, U. 1984, Astrophys. and Space Sci., 104, 337  
 Holmberg, J., & Flynn, C. 2000, MNRAS, in press  
 Kuijken, K., & Gilmore, G. 1989, MNRAS, 239, 605  
 Lépine, J. R. D., & Amaral, L. H. 1995, Mem. Soc. Astron. Ital., 66, 649  
 Merrifield, M. R. 1992, AJ, 103, 1552  
 Nikiforov, I. I. 2000, these proceedings  
 Ninković, S. 1992, Astron. Nachr., 313, 83  
 Ninković, S., & Petrovskaya, I. V. 1992, Sov. Astron., 36, 478  
 Petrovskaya, I. V. 1992, A&A Trans., 3, 87  
 Porcel, C., Battaner, E., & Jiménez-Vicente, J. 1997, A&A, 322, 103