

Initial conditions for massive star birth – Infrared dark clouds

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Abstract. We summarize the properties of Infrared Dark Clouds, massive, dense, and cool aggregations of interstellar gas and dust that are found throughout the Galaxy in projection against the strong mid-infrared background. We describe their distribution and give an overview of their physical properties and chemistry. These objects appear to be the progenitors of high-mass stars and star clusters, but seem to be largely devoid of star formation, which however is taking place in localized spots.

Keywords. stars: formation, ISM: clouds, ISM: molecules, infrared: general, radio lines: ISM, masers

1. Introduction

Observations of dust and molecules provide almost all of the accessible information on deeply embedded high-mass (pre-)protostars, whose emission is frequently not detected even at mid-infrared wavelengths. As evidenced by, a.o., many contributions to this symposium, much current effort is expended on extensive surveys of unequivocal signposts of high-mass star formation, such as sources with tell-tale far-infrared spectral energy distributions and hot, dense molecular cores highlighted by maser emission in the methanol and water molecules.

All the regions traced by these signposts are containing already formed (proto)stars and only recently have clouds with the potential of forming high-mass stars and/or clusters, but still yet largely devoid of stellar objects, been identified: Infrared Dark Clouds, whose observational status we shall summarize here.

2. Infrared Dark Clouds

First recognized in mid-infrared images from the Infrared Space Observatory (ISO) and Midcourse Space Experiment (MSX) satellites Infrared Dark Clouds (IRDCs) appear in silhouette against the Galactic mid-infrared (MIR) background, frequently in filamentary shapes. Perault *et al.* (1996) report that ISOGAL[†] images show “unexpectedly, a number of regions which are optically thick at $15\mu\text{m}$ which are likely due to absorption” and “a convincing correlation with a depletion in $2\mu\text{m}$ source counts”. They estimate $A_V > 25$ and put forward “that these would be good candidates for precursors of star formation sites.” Even before IRDCs became generally known as a distinct class of objects, Lis & Menten (1998) found absorption in the $45\mu\text{m}$ ISO LWS detector band against the MIR background and emission in the $173\mu\text{m}$ band toward M0.25+0.11, the low Galactic longitude end of the Galactic center “dust ridge”, a string of submillimeter (submm) condensations found by Lis & Carlstrom (1994) which terminates with the prominent Sgr B2 star-forming region at its high longitude end. M0.25+0.11 was studied, even

[†] ISOGAL is a 7 and $15\mu\text{m}$ survey with ISOCAM (the ISO 3 – $20\mu\text{m}$ camera) of $\sim 12^\circ$ in the Galactic Plane interior to $|l| = 45^\circ$.

earlier, in detail by Lis *et al.* (1994), who found very little, if any (a weak H₂O maser) signs of star formation in it. Lis & Menten performed grey body spectral fits to the far-infrared data combined with their previous 350 – 800 μm submillimeter measurements and obtained a low temperature, ~ 18 K, for the bulk of the dust in M0.25+0.11's core. In addition, they found that the grain emissivity is a very steep function of frequency ($\nu^{2.8}$; see §4.1). Lis & Menten derived a gigantic mass of $1 \times 10^6 M_{\odot}$ for this object, comparable to the core of the “mini starburst region” Sgr B2.

The first extensive dataset on IRDCs was compiled with the SPIRIT III infrared telescope aboard the Midcourse Space Experiment (MSX; see Price 1995), which surveyed the whole Galactic plane in a $b = \pm 5^{\circ}$ wide strip (Price *et al.* 2001) in four MIR spectral bands between 6 and 25 μm at a spatial resolution of $\sim 18''3$. In an initial census of a $\sim 180^{\circ}$ long strip of the Galactic plane (between $269^{\circ} < l < 91^{\circ}$, $b = \pm 0.5^{\circ}$), Egan *et al.* (1998) find ~ 2000 “compact objects seen in absorption against bright mid-infrared emission from the Galactic plane. Examination of MSX and IRAS images of these objects reveal that they are dark from 7 to 100 μm .”

The IRDCs are best identified in the 8 μm MSX “A” band, because, first, the 7.7 and 8.6 μm PAH features associated with star-forming regions contribute to a brighter background emission and, second, the MSX A band is more sensitive than the satellite's other bands.

M0.25+0.11 and the other condensations of the Galactic center dust ridge also appear conspicuously in absorption on an 8 μm MSX image presented by Egan *et al.* 1998.

Hennebelle *et al.* (2001) in a systematic analysis of the ISO GAL images extracted a total of ~ 450 IRDCs, for which they derive 15 μm opacities of 1 to 4.

The four newly-detected massive and dense cold cores identified by Garay *et al.* (2004; see also his contribution to these proceedings) also represent an interesting sample of IRDCs. These objects are not mid-IR but mm-continuum selected: The 1.2 mm dust emission reveals massive ($M > 400M_{\odot}$) and cold ($T < 16$ K) cores.

How do IRDCs compare to the Orion Molecular Cloud I (OMC-1), probably the best-studied high-mass molecular cloud/star-forming region complex? In Fig. 1, we show the 1.2 mm dust continuum emission from OMC-1 mapped with MAMBO† and the 850 μm SCUBA‡ dust continuum map of IRDC G11.11-0.12 (Johnstone *et al.* 2003). At a distance of 3.6 kpc (see §4.1) G11.11-0.12 has a remarkable resemblance with the integral-shaped Orion filament, both, in structure and in dimensions. The bright MIR emission along the Galactic plane favours the identification of a massive cold cloud as infrared dark, which the OMC-1 region is not due to the absence of a MIR background caused by its location outside of the Galactic plane. One glaring difference though exists between the two maps: the prominent maximum in OMC-1, marking the very active BN/KL high-mass stars-forming region (see §5).

3. Galactic distribution and distances

Using a standard Galactic rotation curve, Carey *et al.* (1998) determined kinematic distances for some of the IRDCs in their sample. They obtain distances between 2.2 and 4.8 kpc, proving without a doubt that the clouds are not local. Their distances agree with those of HII regions in their vicinity.

† The **MA**x-Planck **B**olometer array is operated at the IRAM 30m telescope on Pico Veleta, Spain.

‡ The **S**ubmillimeter **C**ommon **U**ser **B**olometer **A**rray is operated at the 15 m James-Cleck-Maxwell Telescope on Mauna Kea, Hawaii.

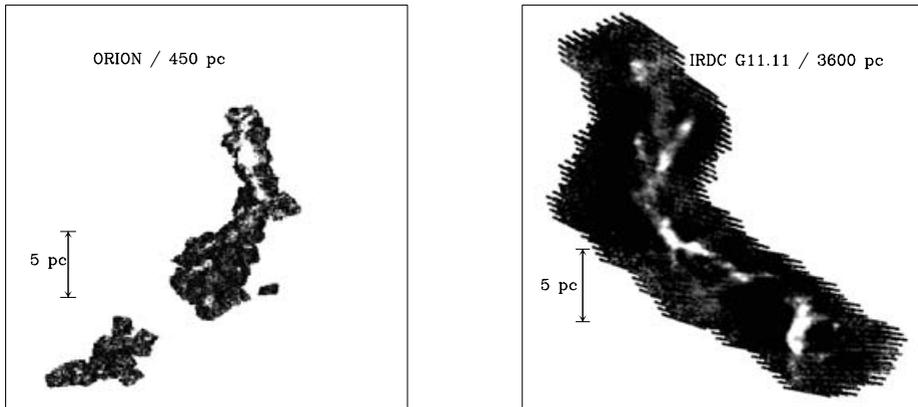


Figure 1. *Left panel:* 1.2 mm dust continuum map of the Orion Molecular Cloud 1 (courtesy T. Stanke). *Right panel:* 850 μm map of G11.11-0.12 (Johnstone *et al.* 2003)

Recently, Simon *et al.* (2004) prepared a catalogue of ~ 380 IRDCs identified from the MSX survey, based on the morphological correlation of MIR extinction and ^{13}CO emission, as observed in the BU-FCRAO Galactic ring survey (GRS) of the inner Milky Way. They find that the majority of the dark clouds are concentrated in the Galactic ring at a galactocentric radius of 5 kpc. The kinematic distances derived range from 2 – 9 kpc. In Table 1, we list the source parameters of all the IRDCs for which data at wavelengths other than the MIR have been collected.

The cloud sizes as reported by Carey *et al.* (1998) are from 0.4 – 15 pc while Teyssier *et al.* (2002) report structures of filaments down to sizes ≤ 1 pc.

4. Physical parameters of IRDCs

4.1. Density, column density and temperature

Based on observations of the formaldehyde molecule (H_2CO), Carey *et al.* (1998) argue that IRDCs objects are dense ($n > 10^5 \text{ cm}^{-3}$), cold ($T < 20 \text{ K}$) cores, apparently without surrounding envelopes. However, later work by Teyssier *et al.* on a different IRDC sample report that large field maps obtained with the 4-m Nanten telescope in the ^{13}CO $J = 1 - 0$ line, which probes relatively low densities, (Zagury *et al.*, unpublished data) indicate that at least these IRDCs may indeed have lower density envelopes. We have confirmed this for the Carey *et al.* sample with hitherto unpublished ^{13}CO $J = 1 - 0$ maps retrieved from the GRS.

Carey *et al.* (1998) conclude that the IRDCs they have studied have extinctions in excess of 2 mag at 8 μm . Using the infrared (visual through 30 μm) extinction law Lutz *et al.* (1996) derived for the Galactic center, this indicates visual extinctions of > 30 mag. How representative Lutz *et al.*'s law is for other lines of sight is unknown. Certainly it does lack the pronounced minimum for standard graphite-silicate mixes in the 4 – 8 μm range predicted by Draine & Lee (1984), which seems well-established by observations toward various lines-of-sight (see references therein).

Actually, IRDC opacities at two wavelengths (e.g. 7 and 15 μm) can be used as a check on the extinction curve and its possible variation with different lines of sight; see Teyssier *et al.* 2002. These authors find a (marginally) lower 7 to 15 μm opacity ratio for clouds located away from the Galactic center compared to clouds that appear in the Galactic center direction. For the latter value they derive ~ 1 , which is consistent with Lutz *et al.*'s law.

Table 1. List of all IRDCs with (sub)mm-wavelength data

Name	R.A.	Decl.	V_{LSR}	d	$N(\text{H} + \text{H}_2)$	T_{kin}
	(J2000.0)	(J2000.0)	(km s^{-1})	(kpc)	(10^{22} cm^{-2})	(K)
DF+04.36-0.06 ^t	17:55:53.07	-25:13:18.7	11.4	3.5	—	—
DF+09.86-0.04 ^t	18:07:37.22	-20:25:54.5	17.7	2.8	6.1	10.0
DF+15.05+0.09 ^t	18:17:37.87	-15:48:59.9	29.9	3.1	12.6	—
DF+18.56-0.15 ^t	18:25:19.52	-12:49:57.0	50.5	4.0	—	8.0
DF+18.79-0.03 ^t	18:25:19.84	-12:34:23.1	—	3.6	—	—
DF+25.90-0.17 ^t	18:39:10.13	-06:19:58.8	—	5.5	—	—
DF+30.23-0.20 ^t	18:47:13.16	-02:29:44.7	104.7	6.7	11.1	8.0
DF+30.31-0.28 ^t	18:47:39.03	-02:27:39.8	—	6.3	—	—
DF+30.36+0.11 ^t	18:46:21.16	-02:14:19.0	96.1	5.9	—	—
DF+30.36-0.27 ^t	18:47:42.37	-02:24:43.2	—	6.9	—	—
DF+31.03+0.27 ^t	18:47:00.39	-01:34:10.0	77.8	4.9	11.1	10.0
DF+36.95+0.22 ^t	18:57:59.51	+03:40:33.3	—	5.0	—	—
DF+51.47+0.00 ^t	19:26:12.74	+16:26:12.6	54.7	5.3	7.7	10.0
G353.85+0.23 P1 ^c	17:29:16.5	-34:00:06	—	—	—	—
G353.51-0.33 P1 ^c	17:30:26.0	-34:41:48	—	—	—	—
G357.51+0.33 P1 ^c	17:40:49.9	-31:14:50	—	—	—	—
G10.74-0.13 P1 ^c	18:09:45.9	-19:42:04	—	—	—	—
G11.11-0.12P1 ^p	18:10:29.27	-19:22:40.3	29.2	3.6	1.7	13.5
G19.30+0.07 ^p	18:25:56.78	-12:04:25.0	26.3	2.2	—	18.5
G24.72-0.75 ^p	18:36:21.07	-07:41:37.7	56.4	3.6	—	20.3
G24.63+0.15 ^p	18:35:40.44	-07:18:42.3	54.2	3.6	—	14.4
G28.34+0.06P1 ^p	18:42:50.9	-04:03:14	78.4	4.8	3.3	16.6
G28.34+0.06P2 ^p	18:42:52.4	-03:59:54	78.4	4.8	9.3	16.0
G33.71-0.01 ^p	18:52:53.81	+00:41:06.4	104.2	7.2	—	17.2
G79.27+0.38 ^p	20:31:59.61	+40:18:26.4	1.2	1.0	8.3	11.7
G79.34+0.33 ^p	20:32:26.20	+40:19:40.9	0.1	1.0	8.8	14.6
G81.50+0.14 ^p	20:40:08.29	+41:56:26.4	8.7	1.3	—	16.6

Notes: Columns are name, right ascension, declination, LSR velocity, distance, total column density, and kinetic temperature. In sources with a distance ambiguity, the near distance was chosen. ^t refers to sources from Teyssier *et al.* 2002, ^c from Carey *et al.* (2000), 'P1/P2' refers to the brightest submm peaks of Carey *et al.* (2000), ^p means that the position of the brightest ammonia peak is given rather than the submm peak position (see Pillai *et al.* 2005a).

Using the relation given by Bohlin *et al.* (1978), assuming that all the hydrogen is in molecular form and a “standard” ratio of total to selective extinction of 3.1, hydrogen column densities in excess of $3 \times 10^{22} \text{ cm}^{-2}$ is derived.

An independent H_2 column density estimate can be obtained from observations of (sub)millimeter dust emission, which, in addition, also allow determination of the cloud (gas+dust) mass (see, e.g., Mezger *et al.* 1987, 1990; Lis & Menten 1998). At (sub)millimeter wavelengths the dust emission is generally optically thin over most of the volume of an interstellar cloud. Thus, at wavelength λ the measured flux density, S_λ , is given by $\int B_\lambda(T_{\text{D}})(1 - e^{-\tau_\lambda})d\Omega$, where T_{D} is the temperature of the dust, τ_λ its optical depth and B_λ is the Planck black body brightness. The integration is either over the beam solid angle for a point-like source or over the source’s angular extent, if the latter is extended. τ_λ is proportional to the hydrogen column density, $N(\text{H}_2)$, and the dust absorption cross section per hydrogen atom σ_λ , which itself is assumed to be proportional to $\lambda^{-\beta}$.

Using all of the above, one finds that the H_2 column density is related to S_λ as $N(\text{H}_2) \propto S_\lambda \lambda^{3+\beta} (e^{hc/\lambda T_{\text{D}}} - 1)$. The cloud mass, M , is proportional to $N(\text{H}_2)D^2$, where D is the cloud’s distance.

Carey *et al.* (2000) imaged a sample of 8 IRDCs in 450 and 850 μm dust emission using SCUBA. Since it is impossible to determine, both, T_{D} and β with just two data points, Carey *et al.* calculated dust color temperatures for three different values of β , 1.5, 1.75, and 2 and note that higher β -values (meaning lower temperatures) yield a better fit to the low temperatures Carey *et al.* (1998) obtained from H_2CO observations. The (high) column densities implied by choosing $\beta = 2$ are around $5 \times 10^{22} \text{ cm}^{-2}$ for 4 sources of their sample and around $13 \times 10^{22} \text{ cm}^{-2}$ for three. Three of the cores corresponding to the brightest submm peaks have masses around $100 M_{\odot}$, two other have 400 and $1200 M_{\odot}$, respectively. Two clouds in the Cygnus region have masses around $40 M_{\odot}$, but we note that distance estimates for that region are very uncertain and probably a short distance (1 kpc used by Carey *et al.* 1998) was chosen for the latter calculations. For $\beta = 1.75$, all these values are to be reduced by a factor of 2.

As reported above (see §2), for M0.25 Lis & Menten derive a very high value of β of 2.8. They take that to indicate the presence of dust grains covered with thick ice mantles.

The values for the high densities and low temperatures deduced by Egan *et al.* (1998) are confirmed by Carey *et al.* (1998), who made, for 10 IRDCs, millimeter-wavelength observations of H_2CO , which is a well-established density probe (Mangum & Wootten 1993; Mundy *et al.* 1987). Since they observed several transitions, they were able, using a Large Velocity Gradient method, to determine temperatures and abundances. Unfortunately, these authors do not discuss their results source by source, but only give general statements.

Leurini in her dissertation (Bonn University; see also Leurini *et al.* 2004) has shown that methanol (CH_3OH) is a highly useful interstellar density and temperature probe. Consequently, she conducted observations of IRDCs in a selected series of lines of that molecule, which she showed to be overabundant in these sources (see §4.3) and, thus, easy to detect. Leurini also corroborates the high densities ($\sim 10^5 - 10^6 \text{ cm}^{-3}$) indicated by the H_2CO data. However, she only observed positions of submillimeter emission peaks, some of which (if not all) harbour embedded sources. Her analysis does, thus, not necessarily apply to general, cool IRDC material, but to gas that is influenced by embedded protostars (see §5). This is reflected most directly in the high-velocity outflows seen in some sources, e.g. G11.11-0.12 and in the high kinetic temperatures of order 40 to 60 K she derives for these.

4.2. Morphology

A significant fraction of IRDCs (although not all) are filamentary. Are IRDCs really filaments, i.e. elongated cylinders, or are they sheets seen edge-on? This is actually an important question linked to their evolutionary state.

Larson (1985), using theoretical arguments and numerical simulations, argues that fragmentation is unlikely to occur in an initially uniform cloud. Either an initial anisotropy or rotation or a magnetic field will in general cause the cloud to collapse toward a flattened or filamentary structure. Once overall collapse has been halted and approximate equilibrium has been established, gravitational instability can cause the resulting sheet or filament to break into fragments of a characteristic mass that depends on the temperature and the surface density of the cloud.

Larson's arguments are supported by the work of Miyama *et al.* (1987a,b), who investigated the fragmentation instability of an isothermal gas layer, in order to see whether the observed structures of many dense interstellar clouds are the results of fragmentation of sheet-like clouds. Linear perturbation theory predicts fragmentation of a parent sheet-like cloud in elongated structures, and using further nonlinear analysis they found

that, if fragments are initially elongated, they become elongated more and more as they go on collapsing, ending up as very slender cylinders, which fragment further.

The G11.11-0.12 IRDC has a distinct filamentary appearance (see Figs. 1 and 4). Using a sophisticated computational technique Fiege *et al.* (2004) compared observations to three different models of self-gravitating, pressure-truncated filaments, namely the non-magnetic Ostriker (1964) model, and two magnetic models from the literature. Analysing the 850 μm SCUBA observations of G11.11-0.12, Johnstone *et al.* (2003) concluded that this source has a much steeper $r^{-\alpha}$, ($\alpha \gtrsim 3$) radial density profile than other (lower mass, lower extinction) filaments, where the density varies approximately as r^{-2} . This steep density profile is consistent with the Ostriker model. After a wider search of parameter space, Fiege *et al.* conclude that the observed radial structure of G11.11-0.12 can be understood in the context of all three models. Discrimination between the different models may be possible with polarization measurements as the magnetic models predict dominant poloidal magnetic fields that are dynamically significant; G11.11-0.12 may be radially supported by a poloidal field. Fiege *et al.* predict polarization patterns expected for both magnetic models, which produce different polarization patterns. Polarimetry should, thus, be able to distinguish between the two magnetic models or a non-magnetic model.

An instrument of choice will be PolKa, the Polarization Kamera designed to be used together with the bolometer arrays developed at the Max-Planck-Institut für Radioastronomie, for example with the 295 element 870 μm Large APEX BOLometer CAmera (LABOCA), which is soon to be installed at the 12m Atacama Pathfinder Experiment telescope (Siringo *et al.* 2004).

To decide the sheet or cylinder question, let us consider G11.11-0.12. That cloud is at a distance of 3.6 kpc (Carey *et al.* 1998). The elongated submillimeter emission has an extent of 24 pc in the long and an average ~ 0.8 pc in the short axis. The density is uncertain: values between 10^5 and 10^6 cm^{-3} have been derived by Carey *et al.* (1998) from their H_2CO data, Johnstone *et al.* (2003) argue for a few times 10^4 cm^{-3} . The H_2 column density is between 0.2 and 2×10^{22} cm^{-2} (Carey *et al.* 2000). Using the extremes of these values, we find that the extent along the line of sight must be between a few times 10^{-3} and 0.5 pc, definitely ruling out a sheet seen edge on.

4.3. Chemistry of IRDCs

Complex organic (i.e. O- and C-bearing) molecules in the interstellar medium are mostly found in hot, dense cores surrounded by high-mass protostars. They frequently have very high over abundances (factors of 100 – 1000) compared to dark cloud values. These are usually explained by the evaporation of icy dust grain mantles on which these molecules are formed in a cooler phase in the clouds lifetime by hydrogenation of CO to H_2CO . Further hydrogenation leads to even more complex species. (Relatively) complex molecules are also found in cold dark clouds, with TMC-1 being a prominent nearby example (See, e.g. Kaifu *et al.* 2004). However, in the latter they all have very small abundances and are observable only because of TMC-1's proximity (yielding a high filling factor) and its moderately high density (10^4 cm^{-3}) leading to substantial beam-averaged column densities. This makes exotic (but not organic in the strict sense of the work) species detectable, such as the polyene carbon chains (Kaifu *et al.* 2004).

What is the organic content of normal molecular clouds? This, essentially, is an unanswered question (the one example TMC-1 aside). Its answer has profound impact on astrochemistry (are grain mantles really needed to form these molecules?) and even astrobiology. Their high column densities make IRDCs ideal laboratories to address this question and potentially detect complex molecules. Such species might be present and

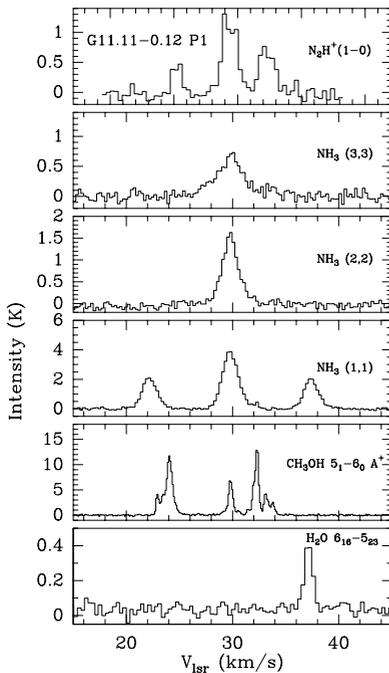


Figure 2. *top to bottom:* Spectra of the N_2H^+ (1-0), NH_3 (J, K) = (3, 3), (2, 2), and (1, 1) lines and the CH_3OH 6.7 GHz and H_2O 22.2 GHz maser transitions. All spectra were taken towards the submm peak position P1 given by Carey *et al.* 2000 (see Table 1).

widespread in many clouds, but would be rendered undetectable because of the modest column densities of ordinary clouds: Lines from almost all molecules significantly rarer than NH_3 or CH_3OH will most likely be optically thin, which makes their line intensity directly proportional to the column density.

Observations of molecules in IRDCs so far have concentrated on just a few species: CO (several isotopomers), H_2CO , CH_3OH , and NH_3 . In addition, Teyssier *et al.* (2002) observed several HC_3N lines and two k -series of CH_3CCH . The latter, a symmetric top, can be used as a temperature probe and its observations yield values for the kinetic temperature, T_{kin} , between 8 and 25 K; the higher values found toward embedded objects. Large Velocity Gradient (LVG) model calculations of HC_3N , ^{13}CO , and C^{18}O yield densities larger than 10^5 cm^{-3} in the densest parts. Teyssier *et al.* ascribe the relatively low observed intensities (a factor of a few lower than in TMC-1) to a very low kinetic temperature (difficult to understand, as TMC-1 is cold, too, ≈ 10 K), a small filling factor or depletion on grains.

In Fig. 2, we show the spectra of different molecules observed towards the brightest submm peak position of G11.11-0.12 (P1).

The NH_3 and CH_3OH observations produced interesting results: CH_3OH and NH_3 are overabundant by factors of 5 – 10 relative to “normal” (= lower density) and less turbulent dark clouds, such as TMC-1 (Leurini, dissertation and 2005, Pillai *et al.* 2005a in prep.). In contrast, H_2CO is *underabundant* by a factor of ~ 50 (Carey *et al.* 1998). Given this situation it is completely unclear which molecules might be detectable and which ones not. Could it be that species with emission sufficiently strong and widespread to be easily detectable have until now been completely missed?

Systematic searches for other molecules will yield a more complete picture of the chemistry of IRDCs, which, while certainly interesting in itself, will also shed light on general formation mechanisms of complex molecules. Moreover, they might help identify new temperature and density tracers and allow studies of (molecule-)selective depletion.

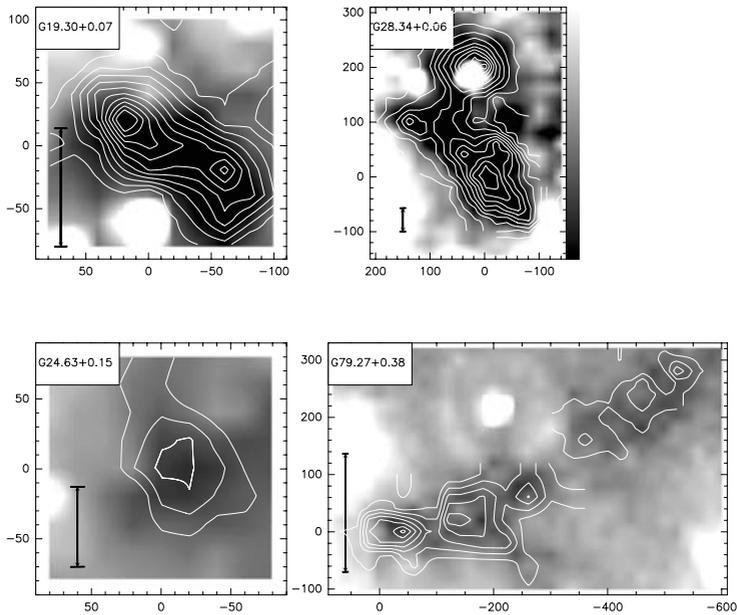


Figure 3. MSX images of the clouds at $8\mu\text{m}$ (greyscale) with NH_3 (1,1) integrated intensity as contours. The contour levels are 2, 4, and 6 times the 1σ noise level. Tick marks are coordinate offsets (in arcseconds) relative to the positions given in Table 1 (from Pillai *et al.* 2005a). The bar marks a distance of 1 pc.

4.3.1. Ammonia

To exploit NH_3 's properties as an excellent molecular cloud thermometer (Danby *et al.* 1988) 10 IRDCs were studied in the course of T. Pillai's dissertation (see also Pillai *et al.* 2005): They were mapped in the $(J, K) = (1, 1)$ and $(2, 2)$ transitions near 1.3 cm wavelength (~ 23.7 GHz) with the MPIfR Effelsberg 100m telescope. The FWHM beam size at the frequencies of the NH_3 lines was $40''$. The NH_3 emission correlates very well with MIR absorption and ammonia peaks (in the following referred to as “cores”) distinctly coincide with dust continuum peaks, as shown in Fig. 3.

Several compact sources were detected within the clouds with sizes smaller than the $\approx 40''$ FWHM beam size. The total gas masses derived for entire clouds from NH_3 data range from $10^3 - 10^4 M_\odot$.

We can constrain the average gas temperatures to $10 \lesssim T \lesssim 20$ K. The temperature distribution within clouds has also been analysed and we find significant temperature gradients, with the temperature rising in outward direction, in all of the cases where we have a good signal-to-noise ratio throughout the map. This outward rise in temperature we find in all except one core can be readily interpreted as influence of the strong external UV field warming up the cloud.

The only case where we find a positive correlation between the gas temperature and the integrated intensity is also the only case where the turbulence seems to increase towards the core. This one has the most evolved core and is also the most massive in our sample. We observe large line widths ($1 \leq \Delta v \leq 3.5$) km s^{-1} , hence turbulence plays an important role in the stability of this IRDC. The column densities translate to extremely high A_V of 55 – 450 mag, therefore early star formation, if any, would be deeply embedded. The virial parameter defined as $\alpha = \frac{5\sigma^2 R}{GM}$ is $1.7 \leq \alpha \leq 4$ for most of the clumps. Hence the cores appear to be unstable against gravitational collapse; in fact

direct evidence for collapse might be revealed from VLA observations we have recently obtained.

The fractional abundance of NH_3 (relative to H_2) is $1 - 6 \times 10^{-8}$. This together with the excellent correlation in morphology of the dust and gas is consistent with the time dependant chemical model for NH_3 of Bergin & Langer (1997) and implies that NH_3 remains undepleted. We can constrain the ages of IRDCs based on this model to $\geq 10^7$ years for H_2 densities $\geq 10^4 \text{ cm}^{-3}$, assuming that the NH_3 has reached its chemical equilibrium abundance. The time scales we derive for the clouds to disperse due to their own internal motions, of a few Myrs, provide a better upper limit to the life time of these clouds. There are significant velocity gradients observed between the cores but we find that they are not attributable to rotation. The effects of external shock/outflow tracers need to be investigated.

Based on the observed line widths, the derived gas temperatures and the NH_3 column densities, we made a comparison of IRDC condensations with objects representing more evolved stages of high-mass star formation like the High-mass Protostellar Objects (HMPOs) studied by Sridharan *et al.* (2002) and Beuther *et al.* (2002). There is a clear trend in temperature from the low temperatures of the IRDCs to typical temperatures of 20 – 30 K for the HMPOs without (significant) HII region to the higher temperatures of UCHII/hot core regions. The line widths in the HMPOs are generally higher than those in the IRDCs.

4.4. Magnetic fields

Theoretical studies suggest that magnetic fields play a crucial role in the star formation process. But contrary to other parameters like density, temperature, the velocity field, and molecular abundances, it has been notoriously difficult to determine B -fields in any regime of the interstellar medium (ISM) from diffuse clouds to dense star-forming cores (see, e.g., Crutcher 1991 and these proceedings).

Virtually the only method for a direct determination of B is the Zeeman effect, which causes a frequency shift of the right-circularly polarized (RCP) relative to the left circularly polarized (LCP) component of a spectral line from a molecule with a suitable electronic structure and also from the 21 cm line of the hydrogen atom.

One of the few interstellar molecules with detectable Zeeman splitting is hydroxyl (OH), whose ground-state hyperfine structure (hfs) transitions near 18 cm wavelength (at 1665, 1667, 1612, and 1720 MHz) have measurable splittings. OH is found in the general molecular interstellar medium and can be detected in clouds with densities \geq a few 10^3 cm^{-3} . However, it is also found, at elevated abundance, in the dense, expanding envelopes of ultracompact HII regions, which have densities $\geq 10^7 \text{ cm}^{-3}$ (Hartquist *et al.* 1995).

B fields of order a few tens of μG have been found from OH Zeeman measurements of low density dark clouds (see, e.g. Goodman *et al.* 1989; Crutcher & Troland 2000), while much stronger, few mG, fields are derived for the much denser maser regions (see, e.g., Fish *et al.* 2003).

While it is relatively easy to measure Zeeman-splitting in OH masers, over the years, large amounts of observing time have been dedicated to measuring Zeeman-splitting in lower density clouds with few, but precious results. A picture has emerged in which the B -field strength increases with density, n , i.e. $|B| \propto n^\kappa$. From theoretical arguments (conservation of magnetic flux and mass) one expects $\kappa = 2/3$ for a collapsing cloud if the B -field is unimportant throughout the collapse and $\kappa = 1/2$ in the opposite case (Crutcher 1991). Models of ambipolar diffusion-driven cloud contraction deliver $\kappa \approx 0.47$ (Fiedler & Mouschovias 1993).

See Fig. 1 of Padoan & Nordlund 1999 for a recent compilation of measured B -field strengths vs. density. It is apparent from this figure that there is a dearth of B -field data points for densities between 10^5 and 10^6 cm^{-3} . B -field measurements of IRDCs will probe just this highly interesting portion of parameter space.

A preliminary survey with the Effelsberg 100m telescope showed OH absorption in both the 1665 and the 1667 MHz hyperfine lines with total (Stokes I) intensities of order -1 K or deeper in several IRDCs. Given the IRDCs' densities cited above we would expect B of order several hundred mG, similar to the values found in high-mass star-forming regions. A 2 hour integration on one source, G yielded an upper limit of 135 μG at a 99% (3σ) confidence level, consistent with the upper limits derived by Crutcher *et al.* (1993) for regions of similar density.

The B -field morphology will be determined from submillimeter polarization observations (see 4.2). Feldman *et al.* (2003) report in an abstract SCUBA polarization observations of MSX IRDCs where they quote very high percentage polarizations ($\sim 6\%$) and find that inferred magnetic field directions are correlated with the cloud structure. There seems to be trend for B to align along the direction of a filament. Bright, compact sources in the filaments are much less polarized, and their inferred B -field directions are perpendicular to the orientation of the filaments.

5. Ongoing star formation in IRDCs

While large volumes of IRDCs appear to be devoid of signposts of ongoing star formation, such as ultracompact HII regions and/or CH_3OH , OH or H_2O masers, isolated centers of high or intermediate star formation are found in many clouds.

Teyssier *et al.* (2002) found that OH and class II CH_3OH masers *et al.* (1995) are associated with positions of (not overly pronounced) peak emission from the column density tracer C^{18}O in the IRDCs DF+9.86-0.04 and DF+30.23-0.20. These are close to dust emission peaks. Since CH_3OH masers are unambiguous tracers of high-mass star formation, we have obtained data on the 6.7 GHz CH_3OH maser transition, towards a sample of ~ 50 dark clouds with a high ($> 25\%$) rate of detections.

Maybe to date the best-studied example of a star-forming core in an IRDC is the 850 μm emission peak in G11.11-0.12 studied in detail by Pillai *et al.* (2005a; in press). Coincident with a compact dust continuum source are both, an H_2O and a CH_3OH maser as shown in the inset of Fig. 4. Interferometric imaging with the Australia Telescope Compact Array show the CH_3OH emission, which has a total velocity spread of ≈ 11 km s^{-1} to have a velocity gradient with emission at different velocities aligned in a line, reminiscent of a disk. Other persuasive arguments for an embedded source are the detection of emission in the high excitation (3,3) line of ammonia with a wider linewidth than the lower excitation (1,1) and (2,2) lines (see Fig. 2). Model fits to all three NH_3 lines indicate a compact source with a size of $\approx 3''$, characterized by a rotation temperature of 60 K, while the more extended emission from the ambient cloud has a rotation temperature of 15 K. The NH_3 column density of the hot, compact component is 9 times higher than that of the cool extended one. Finally, the infrared spectral energy distribution is best modelled by a source with a luminosity of $1200 L_\odot$, corresponding to a ZAMS star of mass $8 M_\odot$. K_s -band 2MASS data show what possibly is reflected light emanating from the protostellar source, which is embedded in a compact mm-wavelength dust continuum source imaged with the Berkeley-Illinois-Maryland Array (BIMA).

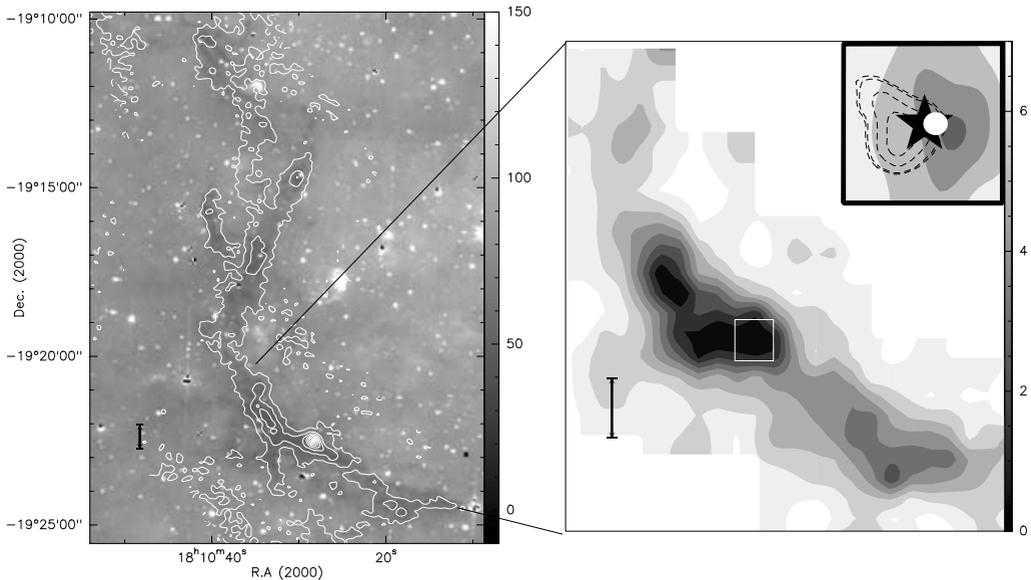


Figure 4. *Left:* The $8\ \mu\text{m}$ image of G11.11-0.12 with the SCUBA $850\ \mu\text{m}$ map (Carey *et al.* 2000) overlaid. *Right:* Southern filament of the SCUBA map in grey scale. The bar marks a distance of 1 pc. The square delineates the position of the active star formation site P1, details of which are shown in the right upper corner inset. Here the greyscale shows a BIMA 3 mm continuum image and a 2MASS K_s band image in shown in dashed contours. The star denotes the H_2O maser position and the filled circle the CH_3OH maser position (from Pillai *et al.* 2005b).

6. Conclusions

The earliest phases of high-mass star formation are expected to be massive (a few hundred to thousand M_\odot), cold (10 – 20 K) and quiescent clouds, emitting primarily at (sub)mm wavelengths and containing no obvious IR sources or star formation tracers. One approach to identify the earliest, cold phases of massive star formation is to search for objects which appear in absorption at MIR wavelengths. Thus IRDCs are the most potential candidates for studying these initial conditions. Centimeter through submm observations reveal that typical IRDCs have gas densities $> 10^6\ \text{cm}^{-3}$, temperatures $< 20\ \text{K}$ and sizes of 1 – 10 pc but studies of their star formation content are still rare.

Studies up to now seem to show that they are not all cold and quiescent. IRDCs appear to harbour sources of different evolutionary stages, not all of them necessarily in the high-mass regime. A better classification scheme based on molecular gas content, MIR contrast and extent is needed to compare IRDCs with local molecular cloud complexes (not clouds). Extensive studies of their energetics, kinematics and chemistry are essential to ascertain their role in forming stars, massive or otherwise. These would be the ideal test grounds for testing the present theories of forming massive stars via turbulent cores (McKee & Tan 2003; see also their contributions to these proceedings). We will need large, Galaxy wide surveys to understand the formation of IRDCs and their lifetimes.

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