

DAILY RADIO FREQUENCY OBSERVATIONS OF SELECTED OBJECTS

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In Table I we present the list of 38 celestial objects that have been observed since January 1978 at 2.7 and 8.1 GHz with the Green Bank interferometer. The sources fall naturally into three categories: radio stars, possibly Galactic sources, and extragalactic sources. SS433, Cyg X-3, and each extragalactic source is measured several times per day while the other sources are measured once every three days. Reports on the entire program will be found in Geldzahler *et al.* (1983a), and on specific sources: SS433--Johnston *et al.* (1983a), BL Lac--Johnston *et al.* (1983b), Cyg X-3--Geldzahler *et al.* (1983b) and elsewhere in this volume), and CTA 26--Spencer *et al.* (1983).

We have defined for the variable sources a "rapidity index" which gives the number of maxima/year. This index includes major outbursts as well as "flickering". We also show in Table I the value of k ($-d(\log S_{\text{max}})/d(\log \lambda)$). The values of this index fall into three groups: $k > 0$, $k \sim 0$, and $k \sim -0.4$. A uniform source that is initially optically thick and whose energy losses occur primarily through adiabatic expansion should yield $k = -1$ (c.f. van der Laan 1966). We find that $k > 0$ when we have an optically thin object such as SS433 or Cyg X-3 during outburst. To make the value of $k \sim -0.4$ more agreeable with the standard model, we suggest the uniformity should be replaced by a variable opacity throughout the source. Finally $k \sim 0$ in those sources, such as the "quiescent" Cyg X-3, where repeated, rapid flickering has stretched and weakened the magnetic field in the immediate vicinity of the source.

REFERENCES

- Geldzahler, B.J. *et al.*: 1983a, to be submitted to A. J.
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Table I. List of Program Sources

Source	ID	Observing Interval	a	Redshift	b	c	d	Other Names
					RI	k		
a. Radio Stars								
0236+610		5001-5007			I			LSI 61 +303
0323+285		4976-5006			--			UX A+1
0334+004		4971-5010			28.1:			HR1099
1617-155		4977-5517			--			Sco X-1
1909+048		4067-5017			5.7	+0.53±0.03		SS433
1956+350		4976-5115			--			Cyg X-1
2050+407		4971-5017			--	-0.04±0.5		Cyg X-3
2259+585		4974-5129			--			GF2259+586
b. Possibly Galactic Objects								
0125+628		4976-5517			4.0			G127.11+0.54
2013+370		4983-5517			3.7			G74.89+1.22
c. Extragalactic Objects								
0224+671	Q	3942-5517			3.6			
0235+164	BL	4971-5148			4.1	-0.42±0.05		DD160
0237-234	Q	3941-5148		2.223	1.5			DD-263, PHL8462
0316+413	C	3942-4066			I			3C84
0336-019	Q	3942-5148		0.852	6.0	-0.01±0.03		CTA26
0355+508	EF	3941-5148			3.1			
0402-382	Q	3941-4066		1.417	8.7:			
0727-115	Q	5060-5148			24.9:	-0.09±0.02		
0742+103	EF	3942-4066, 5010-5059			I			
0851+203	BL	3942-5148		0.306:	3.9	-0.48±0.01		OJ287
0923+392	Q	3942-4066		0.699	I			C39.25, DA267, DK340
0964+658	BL	3942-5517			2.1	-0.04±0.02		
1226+023	Q	3943-5517		0.158	I			3C273, 4C02.32, NRA0400, DN044, DA324
1245-197	Q	3942-5517			0.0			
1328+254	Q	3942-5517		1.055	0.0			3C287, 4C25.43, NRA0424, OP247, DA345
1328+307	Q	3942-5517		0.849	0.0			3C286, 4C30.26, NRA0425, OP348, DA346, CTA60
1502+106	Q	3942-5517		1.833	2.7	+0.08±0.03		DR103
1519-273	Q	3943-4067			11.8:			
1641+399	Q	3941-5517		0.595	0.6			3C345, 4C39.48, NRA0513, OS368, DA420
1749+701	BL	3943-5517			2.4	-0.03±0.04		W1
1901+319	Q	3943-4066			I			3C395
2021+614	Q	5010-5059			I			
2037+511	Q	3943-4066		1.686	I			3C418, 4C51.12, NRA0636
2048+312	Q	4970-5517		3.18	22.3			CL4
2134+004	Q	4011-5517		1.936	I			PHL61, DA553, DX057
2200+420	BL	3942-5517		0.0688	5.7	-0.30±0.09		BL Lac
2251+158	Q	3942-5517		0.859	3.9	-0.12±0.01		3C454.3, 4C15.76, OY185, NRA0701, DA506
2345-167	Q	3942-4066		0.600	20.6:			DZ-176

a As of 1 July 1983

b Emission line redshifts taken from Hewitt and Burbidge (1980)

c Rapidity Index: the number of maxima per year; I = variations exist by the value of RI cannot be determined with reliability, -- = no obvious variation, : = value is uncertain due to short time base

d $k = (d \log \text{peak flux density}) / (d \log \text{frequency})$

e In Cyg X-3, $k \approx 0.5$ during optically thin outbursts and $k \approx -0.4$ during optically thick outbursts