

Ejection of Bi-Helical Fields from the Sun

Axel Brandenburg

Nordita, Blegdamsvej 17, DK-2100 Copenhagen Ø, Denmark

Eric G. Blackman

*Department of Physics & Astronomy, University of Rochester,
Rochester NY 14627, USA*

Abstract. It is argued that much of the observed magnetic helicity losses at the solar surface may represent a reduction of an otherwise more dominant nonlinearity of solar and stellar dynamos. This nonlinearity is proportional to the internal twist (as opposed to writhe) of helical and sigmoidal surface structures.

1. Introduction

Dynamo action is possible both for helical and non-helical turbulence [see, e.g., Meneguzzi et al. (1981) for early numerical work]. However, the reason one is particularly interested in helical dynamos is that there is a well-known mechanism (the $\alpha\Omega$ mechanism) driving cyclic behavior together with latitudinal migration of a large scale field. Both cyclic behavior and latitudinal migration are important features of dynamos in the sun and many late-type stars.

The α -effect is formally introduced via the longitudinally averaged induction equation. This effect characterizes the conversion of toroidal mean field into poloidal via a sequence of events, all of which produce similarly oriented loops that are tilted in a clockwise sense against the toroidal direction in the northern hemisphere, and anti-clockwise in the southern. This effect is the result of the combined action of Coriolis force and non-uniformity of the turbulence, as embodied by the formula (e.g. Rüdiger & Kitchatinov 1993)

$$\alpha_0 \approx -\tau^2 \overline{\mathbf{u}^2} \boldsymbol{\Omega} \cdot \nabla \ln(\overline{\rho \mathbf{u}^2}). \quad (1)$$

This equation, which has been obtained using different approaches (see Moffatt 1978), reflects the fact that the turbulent velocity field has attained *kinetic* helicity. The subscript 0 indicates that the nonlinear feedback is not included.

2. Connection with magnetic helicity

It is first of all important to realize that practically no net magnetic helicity can be generated in the sun. What is possible, however, is a *segregation* of the magnetic field into its positively and negatively helical constituents. Such a segregation can occur either in wavenumber space or in real space. The latter can be accomplished by differential rotation acting on field lines crossing the

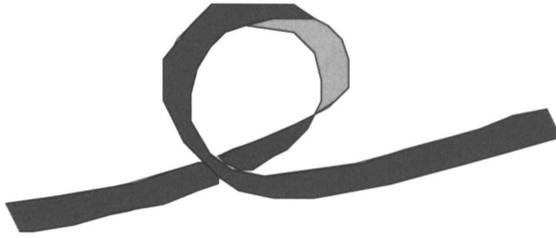


Figure 1. Tilting of the rising tube due to the Coriolis force. Note that the tilting of the rising loop causes also internal twist.

equator, while the former can be the result of the α -effect. This will now be discussed in more detail.

Consider the rise of a toroidal flux tube lifted upward either by thermal or magnetic buoyancy. As it rises, it gets tilted by the Coriolis force and, as a consequence, it must develop some internal twist. This is already quite clear from a simple sketch (Fig. 1), where the flux tube is depicted as a ribbon, so one can trace the induced twist.

Mathematically, magnetic helicity is the sum of writhe and twist helicities, and their sum must stay nearly unchanged (magnetic helicity conservation). In the example shown in Fig. 1, the overall structure of the tube follows the right hand rule, corresponding to positive writhe helicity. At the same time, the interval twist of the tube follows the left hand rule, corresponding to negative twist helicity. On the southern hemisphere both signs would be reversed.

3. Writhe and twist as driver and killer in dynamo theory

We have already alluded to the fact that the tilted magnetic field from the α -effect contributes directly (via many systematic and similar events) to the large scale poloidal field, $\overline{\mathbf{B}}_{\text{pol}}$. Since α is positive, it should give a positive contribution to the current helicity of the mean field, i.e. $\overline{\mathbf{J}} \cdot \overline{\mathbf{B}} > 0$. We also know from helicity spectra of magnetic flux tube experiments that the magnetic helicity from the negative internal twist contributes at scales smaller than from the positive writhe (Blackman & Brandenburg 2003). To put this into numbers, we expect the scales of the large scale poloidal field to be on the order of around 300 Mm, which corresponds to the latitudinal extent of the toroidal flux belts of around 20-30°. The typical scale associated with the twist is expected to be smaller, perhaps 30-300 Mm, but the dividing line between small and large scales may not be very clear. (Obviously, what is small scale to a dynamo theorist may be large scale to a solar observer!)

The internal twist is of tremendous importance in dynamo theory. As explained in the previous section, if the writhe helicity is composed of large scale field, the twist helicity must be composed of small scale field (where ‘small’ could be anywhere between 30-300 Mm!). We may therefore identify the helicity from the internal twist with the current helicity of the small scale field, $\overline{\mathbf{j}} \cdot \overline{\mathbf{b}}$. Here,

lower characters denote the small scale field, i.e. we have decomposed the field as $\mathbf{B} = \overline{\mathbf{B}} + \mathbf{b}$ and likewise the current as $\mathbf{J} = \overline{\mathbf{J}} + \mathbf{j}$.

The significance of the $\overline{\mathbf{j} \cdot \mathbf{b}}$ term is that it gives an extra contribution to the α -effect (Pouquet et al. 1976),

$$\alpha = \alpha_0 + \frac{1}{3} \tau \overline{\mathbf{j} \cdot \mathbf{b}} / \rho, \quad (2)$$

but it acts such as to suppress or quench the dynamo (α_0 and $\overline{\mathbf{j} \cdot \mathbf{b}}$ tend to have opposite signs). Indeed, the $\overline{\mathbf{j} \cdot \mathbf{b}}$ term constitutes the main nonlinearity of α -effect dynamos. There are other nonlinearities, such as a direct suppression of terms in equation (1), but such nonlinearities are never catastrophic in the sense that they do not depend on the magnetic Reynolds number².

4. Why coronal mass ejections might be good for the dynamo

Active regions and coronal mass ejections are the main contributors of magnetic helicity flux from the solar surface (e.g. Démoulin et al. 2002). It might not be appropriate to associate them with small scale losses only. Instead, in view of the combined presence of large and small scale fields in a single field structure (Fig. 1), such losses occur probably simultaneously at large and small scale fields. In this section we argue that this leads to an optimal scenario for dynamos.

Phenomenologically, surface losses of large and small scale helical fields may be described by diffusion terms. This concept has been tested against simulations in the case where losses occur almost entirely at large scales (Brandenburg & Dobler 2001). The results are twofold: the saturation time is formally decreased, but only at the expense of lowering the saturation field strength.

The rest of a possible success story is still speculation, but is based on physical reasoning and a numerical experiment: if losses occur preferentially on small scales, then the saturation amplitude of the large scale field is increased. This is because $\overline{\mathbf{j} \cdot \mathbf{b}}$, the ‘quencher’ in dynamo theory, is constantly being removed. This is confirmed by simulation where small scale magnetic field is artificially removed (Fig. 2). Even in the absence of such losses, the saturation value of the large scale field may be sufficient for the sun, but the cycle period may be too long. The primary importance of small scale losses may actually be to maintain the observed cycle period. This needs further investigation.

5. Conclusions

An important missing link in the story outlined above is, in our view, a simulation of a coronal mass ejection and an analysis in terms of the magnetic helicity budget. At the same time, it will be necessary to improve mean field theory to take helicity fluxes properly into account. For attempts in that direction we refer to recent work by Kleorin and collaborators (2003).

²By contrast, equation (2) can lead to a catastrophic quenching formula: use Keinigs’ (1983) formula for the *saturated* state, $\alpha = -\eta \overline{\mathbf{j} \cdot \mathbf{b}} / \overline{\mathbf{B}^2}$ (which follows from the magnetic helicity equation), eliminate $\overline{\mathbf{j} \cdot \mathbf{b}}$ from (2), to get catastrophic quenching: $\alpha = \alpha_0 / (1 + R_m \overline{\mathbf{B}^2} / B_{\text{eq}}^2)$. Here we have used $R_m = \eta_t / \eta$, with $\eta_t = \frac{1}{3} \tau \overline{\mathbf{u}^2}$ and defined $B_{\text{eq}} = (\mu_0 \rho \overline{\mathbf{u}^2})^{1/2}$ to eliminate τ ; see Eqs (18) and (40) of Blackman & Brandenburg (2002) for a more general formulation.

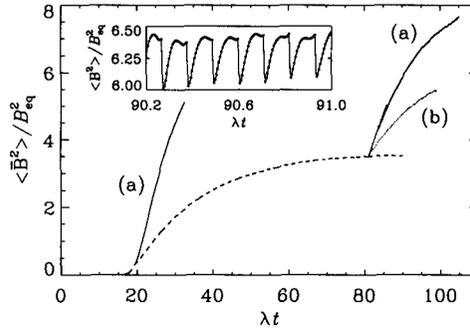


Figure 2. The effect of removing small scale magnetic energy in regular time intervals Δt on the evolution of the large scale field (solid lines). The dashed line gives the evolution of $\langle \overline{B}^2 \rangle$ for Run 3 of Brandenburg (2001), where no such energy removal was included. The two solid lines show the evolution after restarting at $\lambda t = 20$ and $\lambda t = 80$. Time is scaled with the kinematic growth rate λ . The curves labeled (a) give the result for $\Delta t = 0.12\lambda^{-1}$ and those labeled (b) for $\Delta t = 0.4\lambda^{-1}$. The inset shows, for a short time interval, the sudden drop and subsequent recovery of the total (small and large scale) magnetic energy in regular time intervals (adapted from Brandenburg et al. 2002).

References

- Blackman, E. G., & Brandenburg, A. 2002, *ApJ*, 579, 359
 Blackman, E. G., & Brandenburg, A. 2003, *ApJ*, 584, L99
 Brandenburg, A. 2001, *ApJ*, 550, 824
 Brandenburg, A., & Dobler, W. 2001, *A&A*, 369, 329
 Brandenburg, A., Dobler, W., & Subramanian, K. 2002, *AN*, 323, 99
 Démoulin, P., Mandrini, C. H., van Driel-Gesztelyi, L. 2002, *ApJ*, 382, 650
 Keinigs, R. K. 1983, *Phys. Fluids*, 26, 2558
 Kleorin, N. I., Kuzanyan, K., Moss, D., et al. 2003, *A&A*, 409, 1097
 Meneguzzi, M., Frisch, U., & Pouquet, A. 1981, *PRL*, 47, 1060
 Moffatt, H. K. 1978, *Magnetic Field Generation in Electrically Conducting Fluids* (Cambridge University Press, Cambridge)
 Pouquet, A., Frisch, U., & L  orat, J. 1976, *JFM*, 77, 321
 R  diger, G., & Kitchatinov, L. L. 1993, *A&A*, 269, 581