

An observational asteroseismic study of the pulsating B-type stars in the open cluster NGC 884

S. Saesen^{1,2}, M. Briquet^{2,3,†}, C. Aerts^{2,4}, A. Miglio⁵, and F. Carrier²

¹Observatoire de Genève, Université de Genève, Switzerland
email: sophie.saesen@unige.ch

²Instituut voor Sterrenkunde, Katholieke Universiteit Leuven, Belgium

³Institut d'Astrophysique et de Géophysique de l'Université de Liège, Belgium

⁴Department of Astrophysics, Radboud University Nijmegen, The Netherlands

⁵School of Physics and Astronomy, University of Birmingham, UK

Abstract. Recent progress in the seismic interpretation of field β Cep stars has resulted in improvements of the physical description in the stellar structure and evolution model computations of massive stars. Further asteroseismic constraints can be obtained from studying ensembles of stars in a young open cluster, which all have similar age, distance and chemical composition. We present an observational asteroseismic study based on the discovery of numerous multi-periodic and mono-periodic B-type stars in the open cluster NGC 884 (χ Persei). Our study illustrates the current status of ensemble asteroseismology of this young open cluster.

Keywords. open cluster and associations: individual (NGC 884), stars: early-type, stars: oscillations, techniques: photometric

1. Data

We exploited the differential time-resolved multi-colour photometry of a selected field of NGC 884 presented in Saesen *et al.* (2010). The photometry was gathered at 12 sites using 15 different instruments and resulted in almost 77 500 CCD images and 92 hours of photo-electric data. The data span three observation seasons (2005 – 2007) and the resulting precision of the light curves of the brightest stars is 5.7 mmag in V , 6.9 mmag in B , 5.0 mmag in I and 5.3 mmag in U . We identified 75 periodic B-type stars through an automated frequency analysis of the V -filter data and focus on these stars in the current study. We derived effective temperature and luminosity values from absolute photometry in the seven Geneva filters and we combined them with literature values to get the position of the B-type stars in the H-R diagram.

2. Derivation of the pulsational properties of the B-type stars

We performed a detailed frequency analysis with PERIOD04 (Lenz & Breger 2005). We used a weighted and non-weighted frequency analysis for the V light curve, and only a weighted analysis for B , I and U . We considered a peak significant if it has a signal-to-noise ratio above 4. A classical prewhitening scheme was adopted to determine all significant frequencies present in a certain star. As a result, we keep 65 periodic B-type

† Postdoctoral Researcher of the Fund for Scientific Research, Fonds de la Recherche Scientifique – FNRS, Belgium

stars of which 36 are mono-periodic, 16 are bi-periodic, 10 are tri-periodic, 2 are quadrupole-periodic and one star has 9 independent frequencies. The results of the multi-frequency fits in the different colours are presented in Table 4 of Saesen *et al.* (2013).

Since none of the detected modes shows any phase differences in the different filters, we restricted for the mode identification to the photometric amplitude ratios to determine the degree l , following the method of Dupret *et al.* (2003). For each pulsator, we selected stellar equilibrium models that fitted the observed position in the H-R diagram within a 2σ -error box. We then selected for each model and each degree l between 0 and 4 the theoretical frequencies that fit the observed ones, taking into account rotational splitting. We confronted the theoretical amplitude ratios with those observed and made a mode identification by eliminating the degrees for which the theoretical amplitude ratios do not match the observed ones. We securely identified the degree for 12 of the 114 detected frequencies.

3. Relations between pulsation and basic stellar parameters

For the following analyses, we exclude three stars: one ellipsoidal binary and two non-member δ Sct stars. All other 62 B-type stars can be considered members of NGC 884 based on photometric diagrams and membership studies in the literature. We did not find a correlation between the projected rotational velocity and the amplitude or frequency of the modes. The amplitudes of the oscillations, however, decrease as the frequency values increase. A comparison between the observed and theoretical frequency-radius relation allowed us to identify eight β Cep stars in our sample, namely Oo 2246, Oo 2299, Oo 2444, Oo 2488, Oo 2520, Oo 2572, Oo 2601 and Oo 2694.

4. Age estimate of the cluster

The seismic properties of the selected β Cep stars were confronted with those predicted for a dense full grid of stellar models (Briquet *et al.* 2011). We searched for a consistent cluster solution by requesting common cluster parameters, i.e., equal age and initial chemical composition, without specifying their values. We eliminated models based on the observed position of the stars in the H-R diagram and the observed frequency spectrum, allowing for rotational splitting in the first-order approximation. Imposing the identified degrees and measured frequencies of the radial, dipole and quadrupole modes of five β Cep stars led to a seismic cluster age estimate of $\log(\text{age}/\text{yr}) = 7.12 - 7.28$. This is fully compatible with the age estimate obtained from modeling of an eclipsing binary in the cluster ($\log(\text{age}/\text{yr}) > 7.10$, Southworth *et al.* 2004) and illustrates the valuable alternative that cluster asteroseismology can offer for age determinations.

5. Further details

For more information, we refer the interested reader to Saesen *et al.* (2010, 2013).

References

- Briquet, M., Aerts, C., Baglin, A. *et al.* 2011, *A&A*, 527, A112
Dupret, M.-A., De Ridder, J., De Cat, P. *et al.* 2003, *A&A*, 398, 677
Lenz, P. & Breger, M. 2005, *CoAst*, 146, 53
Saesen, S., Carrier, F., Pigulski, A., *et al.* 2010, *A&A*, 515, A16
Saesen, S., Briquet, M., Aerts, C., *et al.* 2013, *AJ*, 146, 102
Southworth, J., Zucker, S., Maxted, P. F. L., & Smalley, B. 2004, *MNRAS*, 355, 986