

Nobeyama 45m CO Galactic Plane Survey: Filament properties and star formation in M17

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Abstract. We present the $^{12}\text{CO}(J=1-0)$, $^{13}\text{CO}(J=1-0)$, and $\text{C}^{18}\text{O}(J=1-0)$ maps of the M17 giant molecular clouds (GMCs) obtained as a part of the Nobeyama 45m CO Galactic Plane Survey. The observations cover the entire area of M17 SW and M17 N clouds at an angular resolution of $\sim 15''$ which corresponds to ~ 0.15 pc. We found that the N cloud consists of a couple of twisted filaments, they are extended in parallel toward the H II region. The typical width of the filaments is ~ 0.5 pc in ^{13}CO intensity map. Most of young stellar objects (YSOs) are located on the filaments which have a bright rim structure in $8\mu\text{m}$ at the filament edge facing the H II region. Furthermore, the time scale of the YSOs formation on the bright rim is comparable with that of NGC 6618 cluster which provides UV photons for the region. This fact indicates that the cluster triggered to form YSOs in N cloud. We also investigated the geometry of the H II region and GMCs by comparing spatial distribution of ^{12}CO velocity channel map and infrared dark cloud, and then found that NGC 6618 is possibly formed by the cloud cloud collision.

Keywords. ISM: clouds, ISM: individual (M17), radio lines: ISM, stars: formation

1. Introduction

Most stars are formed in giant molecular clouds (GMCs) in the Galaxy. Recently, *Herschel* observations reveal the ubiquitous existence of the filamentary structure (e.g., André *et al.* 2010) and demonstrate that spatially complete high-resolution observations are essential to understand the whole picture of star formation in the GMCs. Molecular line observations have an advantage in GMC studies because it provides velocity information (e.g., Nishimura *et al.* 2015) while its limited sensitivity prevent us to obtain large scale GMC images with sub-parsec resolution. Nobeyama 45m CO Galactic Plane Survey (Umemoto *et al.* in prep.) is realized by a newly developed multi-beam receiver FOREST (Minamidani *et al.* in prep.) and the survey provides the most powerful dataset for this purpose.

M17 is one of the most famous H II region excited by the high-mass star cluster NGC 6618 (the region is reviewed by Chini & Hoffmeister 2008). The cluster consists of several O stars and estimated to be very young (~ 0.5 Myr; Hoffmeister *et al.* 2008). The region mainly consists of two GMCs, M17 SW and M17 N (Lada 1976). We observed entire region of the clouds by

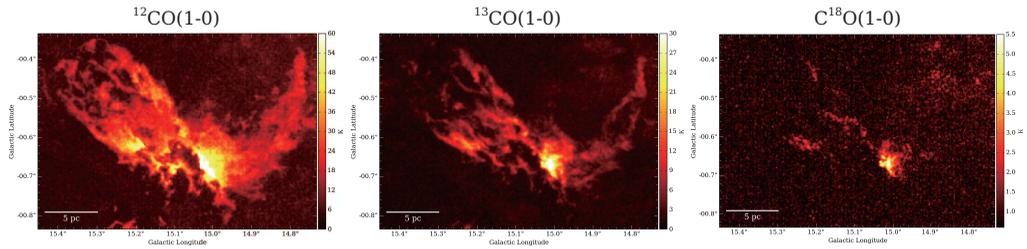


Figure 1. Peak intensity maps of $^{12}\text{CO}(J=1-0)$, $^{13}\text{CO}(J=1-0)$, and $\text{C}^{18}\text{O}(J=1-0)$.

using Nobeyama 45 m telescope and obtained $^{12}\text{CO}(J=1-0)$, $^{13}\text{CO}(J=1-0)$, and $\text{C}^{18}\text{O}(J=1-0)$ images at an angular resolution of $\sim 15''$.

2. Results and discussions

Figure 1 shows the peak intensity maps of CO lines observed with Nobeyama 45 m telescope. The M17 molecular cloud complex is fully covered with improved angular resolution compared with previous $^{12}\text{CO}(J=2-1)$, and $^{13}\text{CO}(J=2-1)$ observation carried out by Povich *et al.* (2009).

We use DisPerSE algorithm (Sousbie 2011) to identify filamentary structures from the cube data of ^{13}CO emission. After the filtering of doubtful candidates, we found that the typical filament width is ~ 0.5 pc and the density profiles seem to be follow Plummer-like function for both SW and N clouds. In the N cloud, almost all young stellar objects (YSOs) identified by Povich *et al.* (2013) are located on the filaments. Most of the YSOs are classified as Class 0/I and a few are Class II/III. All filaments associated with YSOs have bright rim features at $8\mu\text{m}$ with the direction of the cluster indicates that the star formation in N cloud is triggered by the feedback of the cluster.

We compared $^{12}\text{CO}(J=1-0)$ velocity channel maps with infrared images obtained by *Spitzer/IRAC*. We found that 20 km s^{-1} components (blue) are mainly corresponding to dark lane observed in N cloud, while 23 km s^{-1} components (red) have no dark cloud counterparts. Therefore, N cloud filaments (blue component) are located at the near side of the H II region and the red component are located at the far side of the H II region. Furthermore, both blue and red components are seen around the cluster. One possibility which satisfy the results is that the blue and red components were collided and this cloud cloud collision triggered NGC 6618 formation.

References

- André, P., Men'shchikov, A., Bontemps, S., *et al.* 2010, *A&A*, 518, L102
 Chini, R., & Hoffmeister, V. 2008, *Handbook of Star Forming Regions*, Volume II, 625
 Hoffmeister, V. H., Chini, R., Scheyda, C. M., *et al.* 2008, *ApJ*, 686, 310
 Lada, C. J. 1976, *ApJS*, 32, 603
 Minamidani, T., Nishimura, A., Miyamoto, Y., *et al.* in prep.
 Nishimura, A., Tokuda, K., Kimura, K., *et al.* 2015, *ApJS*, 216, 18
 Povich, M. S., Churchwell, E., Bieging, J. H., *et al.* 2009, *ApJ*, 696, 1278
 Povich, M. S., Kuhn, M. A., Getman, K. V., *et al.* 2013, *ApJS*, 209, 31
 Sousbie, T. 2011, *MNRAS*, 414, 350
 Umemoto, T., Minamidani, T., Nishimura, A., *et al.* in prep.