

# Where methanol masers spring

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**Abstract.** We are carrying out a program to observe the 6.7 GHz methanol maser emission in high-mass star-forming regions using large FOV ( $\sim 2'$ ), astrometric, VLBI data. Here we report on the first results of the inner few arc seconds in Cepheus A East. We find a maser filament extending over  $\sim 1.7''$  (1200 AU), straddling the waist of Cep A HW2. The region in which the CH<sub>3</sub>OH masers are found contains several YSO's and it is not clear whether the CH<sub>3</sub>OH masers are associated with several different objects or rather the larger scale surrounding environment.

**Keywords.** masers, stars: formation, stars: individual (Cepheus A)

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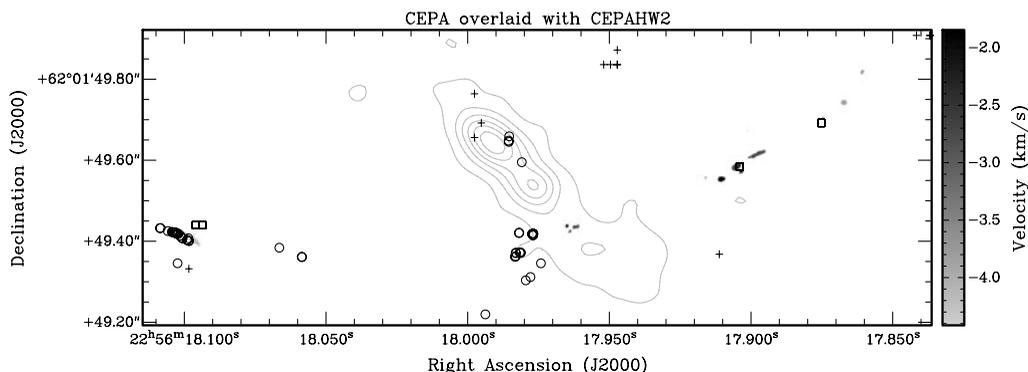
## 1. Introduction

At a distance of 725 pc, Cepheus A East is a nearby complex high-mass star-forming region. At large scale several outflows are seen in many molecular species (e.g. Codella *et al.* 2005). In the centimetre regime the continuum emission is dominated by several compact sources and the bipolar jets associated with Cep A HW2. It is not yet clear how many of the compact sources are associated with individual YSO's. Also, in the region close to HW2 several YSO's have been reported in a number of tracers (Martín-Pintado *et al.* 2005, Curiel *et al.* 2002, Patel *et al.* 2005, Brogan *et al.* 2007). In this poster we report on recent results of the 6.7 GHz methanol masers in the region of Cep A HW2.

## 2. Observations

The observations were carried out in November 2004 with 8 telescopes of the EVN. Cep A was observed for  $\sim 2$ h including calibrators. The observations were split into two 1h blocks, separated by a few hours to improve the uv-coverage. The experiment was setup at a centre frequency of 6668.5142 MHz (rest frequency) with 2 MHz bandwidth and 1024 channels, resulting in a velocity resolution of  $\sim 0.1$  km s<sup>-1</sup>. Both RCP and LCP were recorded. The data was correlated at JIVE with an integration time of 0.25 s.

All calibration and imaging was done in AIPS. Amplitude and bandpass calibration was done on 3C345 and DA193, observed at the beginning and end of each block. For phase calibration and initial self-calibration J2302+6405 was observed every three minutes. The brightest maser feature was then used to determine the absolute position and as a source for further self-calibration, resulting in an rms of  $\sim 3.5$  mJy beam<sup>-1</sup> in the emission free channels.



**Figure 1.** The 6.7 GHz  $\text{CH}_3\text{OH}$  masers (grey scale) overlaid on K band continuum (contours) (Torrelles *et al.* 1998). Also shown are the 12.2 GHz  $\text{CH}_3\text{OH}$  masers (boxes) (Minier *et al.* 2001), water masers (circles) (Vlemmings *et al.* 2006) and OH masers (pluses) (Bartkiewicz *et al.* 2005).

### 3. Discussion

In Cep A we find the  $\text{CH}_3\text{OH}$  maser emission to be arising from an extended  $\sim 1.7''$  ( $\sim 1200$  AU) narrow filamentary structure, which seems to suggest a single underlying structure. Across the whole structure there is a small velocity gradient but without obvious signs of rotation. The maser positions to the west correspond well to the positions of the 12.2 GHz  $\text{CH}_3\text{OH}$  masers measured by Minier *et al.* (2001). To the east the positional agreement is poorer, this could simply be an effect of observations at different epochs. Overall, the velocity of the 6.7 GHz masers agrees well with that of the 12.2 GHz masers.

In all regions except one the  $\text{CH}_3\text{OH}$  masers are spatially offset from other maser species in the region (water and OH). Interestingly, the methanol masers to the east are spatially coincident with the water masers observed by Vlemmings *et al.* (2006), although they show a velocity offset of  $\sim 10 \text{ km s}^{-1}$ . The weak maser emission arising from the very waist of HW2 is very interesting as it may point toward the very object responsible for the outflows/jets. It is however not coincident with the 7 mm continuum object identified by Curiel *et al.* (2006), but a position further to the SW in the jet.

In fact, none of the proposed YSO's in this region (Martín-Pintado *et al.* 2005, Curiel *et al.* 2002, Patel *et al.* 2005, Brogan *et al.* 2007) show a direct, small scale, association with any  $\text{CH}_3\text{OH}$  maser. If the masers are arising in a single structure it seems difficult to attribute their origin to a single object. It is remarkable that a structure encompassing several young stars has such a small velocity gradient.

### References

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