

Bp stars in Orion OB1 association

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Abstract. A total of 85 CP stars of various types are identified among 814 members of the Orion OB1 association. We selected 59 Bp stars, which account for 13.4% of the total number of B type stars in the association. The fraction of peculiar B type stars in the association is found to be twice higher than that of peculiar A type stars.

Magnetic field are found in 22 stars, 17 of them are objects with anomalous helium lines. No significant differences are found between the field strengths in the Bp type stars of the association and Bp type field stars. We identified 17 binaries, which make up 20% of the total number of peculiar stars studied which is the standard ratio for CP stars.

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1. Introduction

The Orion constellation hosts one of the most popular groups of early-type stars in the solar neighborhood - the Orion OB1 association. Blaau (1964) identified four regions inside the association - subgroups: a (corresponds to the northern part), b (the Orion's Belt), c (the region located south of Orion's Belt), and d (the very compact region located in the central part of the association) that slightly differ in age and stellar composition.

Most of the objects in the Orion OB1 association are normal hot main sequence stars; however, the association also includes pre-MS objects, like HAEBE stars, T Tau-type stars, and various anomalous (peculiar) stars. Chemically peculiar (CP) stars differ from normal stars by their anomalous chemical composition which shows up in enhanced or weakened intensity of lines of certain elements in the stellar spectrum.

Renson and Manfroid (2009) published the most detailed catalog of CP stars, which includes more than 8200 objects. Over the past quarter-century many new observations of CP stars have been performed. Our aim is to thoroughly analyze massive chemically peculiar and magnetic stars in the Ori OB1 association using all available data. For a review and analysis of the main studies on this subject, see Romanyuk & Yakunin (2012) and Romanyuk *et al.*, (2013).

2. CP stars in the association

Groups of hot stars in the Ori OB1 association have repeatedly attracted the researcher's attention. Here we consider only the issues related to chemically peculiar stars and magnetic field of these objects.

Borra & Landstreet (1979) discovered very strong magnetic fields in a group of B-type stars with enhanced helium lines in young clusters in Orion. Klochkova (1985) performed spectroscopic observations of 24 CP stars to determine the distance moduli and ages of

Table 1. Age of subgroups and number of normal and CP stars

subgroup	age, log t	all stars	CP stars	fraction, %
Ori OB1a	7.05	311	24	7.7
Ori OB1b	6.23	139	21	15.1
Ori OB1c	6.66	350	37	10.6
Ori OB1d	6.0	14	3	21.4

Table 2. Number of CP stars in different subgroups

Peculiarity type	total	Ori OB1a	Ori OB1b	Ori OB1c	Ori OB1d
Am	23	6	4	13	0
He-strong	7	1	3	1	2
He-weak	27	7	8	12	0
Si, Si+	19	6	4	8	0
Other	9	3	1	4	1

subgroups. Brown *et al.* (1994) reported the results of photometric observations in the Walraven system for 814 stars, for all identified of suspected association members. They determined the effective temperatures, surface gravities, luminosities and masses for all 814 stars. They also determined the distance moduli and showed that the near and far edges of clouds in the Orion OB1 association are located at the distances of about 320 and 500 pc, respectively. We decided to identify chemically peculiar stars among the 814 objects of this list. We consider a star to be peculiar if it appears in the catalog by Renson and Manfroid (2009).

We selected 85 CP stars in the direction of the Ori OB1 association. We list these objects in paper by Romanyuk *et al.*, (2013). Most of them (59 objects) are Bp stars, however, we also found 23 Am and 3 Ap stars. We performed magnetic field observations using the 6m Russian telescope.

The age of subgroups, number of normal and CP stars in each subgroup and fraction of CP stars are presented in Table 1.

We determine distances for most of the stars closer than 250 pc from their Hipparcos parallaxes and we estimate the distances to more distant objects from their temperatures and luminosities. Proper motions have been measured for all CP stars.

The sample of peculiar stars is offset relative to the entire sample both in terms of temperature and luminosity. The fraction of hot stars is greater among peculiar stars. The effective-temperature distribution for the entire sample and CP stars peak at $\log T_e = 3.95$ and $\log T_e = 4.15$ (for details see Romanyuk *et al.*, (2013). The distribution of CP stars in different subgroups is presented in Table 2.

The catalog of CP stars by Renson and Manfroid (2009) includes 23 Am stars in the directions of the association. This is surprising given that low-mass Am stars should not have yet evolved enough to settle onto the main sequence. We therefore decided to verify whether the Am stars in question are foreground objects and not members of the association. Parallaxes are available for 14 of the 23 Am stars and they support conclusively the above hypothesis - these objects are located closer than 300 pc. The distances of remaining nine stars can be determined only from analysis of their temperatures and luminosities. The result of analysis is the same: the distances to Am stars are closer than 300 pc.

Occurrence frequency of CP stars

The fraction of CP stars can be seen to be smallest (7.7%) in the oldest subgroup (a) of the association. It is twice higher (15.1%) in the substantially younger subgroup (b). The fraction of peculiar stars is even higher in the youngest subgroup (d), however, it

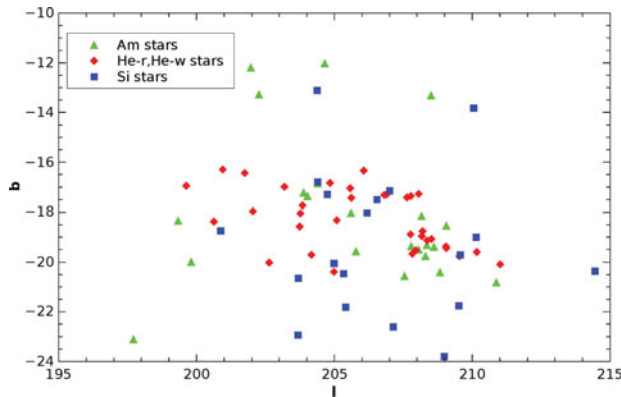


Figure 1. Spatial distribution of CP stars in the Ori OB1 association.

contains too few (only 14) objects to allow any statistical conclusions. The lists of Brown *et al.* (1994) contains 661 stars in the two older groups (a) and (c): 81.2% of the total number of stars in the association. We adopt this fraction as standard. Bp stars are much less concentrated in old subgroups. Thus the fraction of He-weak stars in the old subgroups (a) and (c) is equal to 70%; that of Si-stars, to 52.6%, and that of He-strong stars, to 28.6%. The discrepancies are quite substantial and significant. The predominance of stars with enhanced helium lines is especially conspicuous in the young subgroups (b) and (d). It is interesting that the fraction of He-strong stars in subgroups (b) and (d), whose ages do not exceed 2 Myr, is three times higher than that of He-weak stars in the same subgroups.

Spatial distribution of CP stars.

Spatial distribution of CP stars in the Orion OB1 association is demonstrated on Fig. 1.

The group of Am stars is located separately from the other objects and closer to us. It can be concluded that Am stars are foreground objects which are merely seen projected against the association. Stars with anomalous helium lines belong to cluster and mostly concentrate in subgroups (b) and (c) – 20 objects, whereas subgroup (a) contains only 8 objects. Silicon stars concentrate mostly in subgroups (a) and (c) – 14 out of 18 objects.

Binary CP stars.

We identified 17 physical binaries among 85 CP stars of the association, which makes up a typical fraction of 20%. These stars are distributed by peculiarity type as follows: He-strong – 5 (out of 7), He-weak – 4 (out of 27), SiSi+ – 2 (out of 10), Am – 1 (out of 23), and other types – 5 (out of 9).

The largest fraction of binaries is found among He-strong stars. The two stars with no companion found are HD 36982 and HD 37776. Weak Am stars are very poorly studied, and abnormally low number of binaries may be result of absence of radial velocity measurements. He-weak and SiSi+ stars show the fraction of binaries lower than normal.

3. Magnetic fields

We found 22 magnetic stars in the association of which 21 are Bp stars. Table 3 lists the stars with reliably detected magnetic field.

Eight of the magnetic stars are binary (36.4%). The overwhelming majority (17 out of 22, or 77%) are objects with anomalous helium lines. The fraction of magnetic stars in the inner subgroup (b) is twice higher than in the outer subgroups (a) and (c).

Table 3. Magnetic stars in the Ori OB1 association

Star	Sp Pec	B_e extrema	Star	Sp Pec	B_e extrema
HD 35008	B8 Si	-340	HD 36668	B7 He-wk, Si	-2200/+2000
HD 35298	B6 He-wk	-3000/+3000	HD 36916	B8 He-wk, Si	-1100/0
HD 35456	B7 He-wk	-400/+1080	HD 36955	A2 CrEu	-1300/-410
HD 35502	B6 SrCrSi	-2250/-180	HD 37017	B2 He-strong	-2300/-300
HD 35730	B7 He-wk	-450/+250	HD 37058	B2 He-wk, Sr	-1200/+1200
HD 36313	B8 He-wk	-1500/-1100	HD 37140	B8 SiSr	-1050/+400
HD 36429	B6 He-wk	-840/+160	HD 37479	B2 He-strong	-1600/+3500
HD 36485	B2 He-strong	-3700/+3000	HD 37642	B9 He-wk, Si	-3000/+3000
HD 36526	B8 He-wk, Si	-3500/+3400	HD 37687	B7 He-wk	-600/+500
HD 36540	B7 He-wk	-900/+1030	HD 37776	B2 He-strong	-2000/+2000
HD 36629	B3 He-wk	-1300/+1100	HD 290665	B9 SrCrEu	-1600/+5000

We see no significant differences between magnetic stars of the Ori OB1 association and Bp stars in general in terms of magnetic field strength. However, despite poor statistics, He-strong stars can be seen to possess, on the whole, a factor of 1.5 –2 stronger fields than He-wk stars.

4. Conclusion

We thus identified 85 CP stars in the direction toward the young Ori OB1 association. Our CP stars are distributed by peculiarity types as follows: 23 Am stars, 3 Ap stars and 59 Bp stars. The fraction of peculiar B-type stars in the association is twice higher than the corresponding fraction of peculiar A-type stars. The association includes 22 magnetic stars; 21 of them are Bp stars, and only one is an Ap star. We suggest that when the stars were born in the Orion OB1 association the magnetic fields formed mostly in the objects that later developed helium rather than silicon anomalies.

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